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# The Vera C. Rubin Observatory Data Preview 2

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## ABSTRACT

We present an early version of Rubin Data Preview 2 (DP2), the second data preview from the National Science Foundation–Department of Energy Vera C. Rubin Observatory. Early Data Preview 2 (EDP2) comprises coadds, detection catalogs, and ancillary data products; eventually, the full DP2 will also include single-epoch images and difference images, but any analysis begun on EDP2 will require no modification. DP2 is derived from observations acquired by the LSST Science Camera (LSSTCam) on the Simonyi Survey Telescope at the Summit Facility on Cerro Pachón, Chile, primarily during the on-sky commissioning campaign between 2025-04-25 and 2025-09-21, supplemented by

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observations taken between 2025-10-25 and 2026-01-06 that overlap the commissioning footprint. The DP2 footprint comprises the Science Validation wide-area survey, five Deep Drilling Fields, and a number of targeted small-field regions, including Trifid and Lagoon, Prawn, M49, and New Horizons, all observed as part of the Rubin First Look campaign. Each field was imaged in up to six broad photometric bands, *ugrizy*, and coadded to produce deep imaging covering an estimated 3,000 square degrees. The addition of single-visit-only areas expands the total DP2 footprint to an estimated 15,000 square degrees, with coverage in at least one filter. The median per-visit PSF FWHM across the wide-area survey ranges from 1.17 arcsec in the $z$ band to 1.26 arcsec in $g$ and $r$ bands. COSMOS, the deepest field, reaches estimated coadded $5\sigma$ depths of $u$=26 mag, $g$=26.8 mag, $r$=26.3 mag, $i$=26.1 mag, $z$=25.3 mag, $y$=23.9 mag. DP2 contains 790 million distinct astrophysical objects, of which 299819 solar system objects. DP2 is approximately 59.15 TB in size and is available to Vera C. Rubin Observatory data rights holders via the Rubin Science Platform, a cloud-based environment for the analysis of petascale astronomical data. Based on a roughly five-month primary observing baseline and covering only part of the eventual LSST footprint, DP2's area, depth, and multiband coverage nonetheless support a broad range of early science investigations ahead of LSST Data Release 1.

*Keywords:* Rubin Observatory - LSST



## 1. INTRODUCTION

The National Science Foundation (NSF)–Department of Energy (DOE) Vera C. Rubin Observatory is a ground-based, wide-field optical-near-infrared facility on Cerro Pachón in northern Chile, built to carry out the Legacy Survey of Space and Time (LSST) (Ivezić et al. 2019a). The observatory is named in honor of Vera C. Rubin, whose observations in the latter half of the 20th century established the first compelling empirical case for dark matter (Rubin & Ford 1970; Rubin et al. 1980). Its 10-yr survey will image the entire visible southern sky every 3–4 nights, reaching a coadded depth of $\sim$ 27 mag in the $r$-band across an 18000 square degrees footprint (Bianco et al. 2022).

At the heart of the facility is the 3.2 gigapixels-gigapixel LSST Science Camera (LSSTCam), mounted on the 8.4 m (6.5 m effective aperture) Simonyi Survey Telescope (Stalder et al. 2024). Together they deliver an instantaneous field of view of 9.6 deg$^2$ in six broadband filters, *ugrizy* (320–1050 nm), with seeing-limited image quality. The system's étendue of 319 deg$^2$ m$^2$ exceeds that of any previous optical facility by over an order of magnitude, enabling the rapid, deep, all-sky cadence that underpins Rubin's four core science themes: the nature of dark energy and dark matter; a census of the solar system; the transient and variable optical sky; and the structure of the Milky Way (Ivezić et al. 2019a). An automated Data Management System processes tens of terabytes of images per night to produce science-ready data products within minutes of each observation for a global community of scientists.

This paper presents Rubin's third Data Preview, Data Preview 2 (DP2) (NSF-DOE Vera C. Rubin Observatory 2025), and the first to be based entirely on LSSTCam data. It marks the conclusion of the Rubin commissioning program and the transition to full survey operations. DP2 is a reprocessing of all science-grade exposures acquired with LSSTCam during on-sky commissioning and the initial phase of early operations, prior to the formal start of the LSST on 2026-06-29. As such, it represents a natural boundary and provides a comprehensive dataset that captures the on-sky performance of the as-built system at the end of the Rubin construction project.

Two earlier Data Previews preceded DP2 as part of the Rubin Early Science Program (Guy et al. 2026), established to provide the community with early access to data products and services needed to produce high-impact science during the period spanning commissioning through the first Data Release. Data Preview 0 (DP0) was based on simulated LSST-like data (LSST Dark Energy Science Collaboration (LSST DESC) et al. 2021) and served primarily to stand up early versions of the Rubin Science Platform (RSP) and prepare the community for data access at scale. Data Preview 1 (DP1) (NSF-DOE Vera C. Rubin Observatory 2025) delivered the first data products derived from on-sky observations using science-grade exposures acquired with Rubin Commissioning Camera (LSSTComCam) (Stalder et al. 2022, 2020; Howard et al. 2018; SLAC National Accelerator Laboratory & NSF-DOE Vera C. Rubin Observatory 2024); a 144-megapixel, single-raft instrument used during early commissioning prior to the installation of LSSTCam on the Simonyi Survey Telescope. DP2 supersedes both in scope, fidelity, and scientific depth.

DP2 contains 29756 science-grade exposures acquired over 158 nights between 2025-04-16 and 2025-09-21, covering a total sky area of 15,000 square degrees across 85 fields. The data products include raw and calibrated single-epoch images, coadded images, difference images, source detection catalogs, and a range of derived data products. The release totals approximately 59.15 TB and contains around 790 million distinct astronomical objects. Full release documentation is available at https://dp2.lsst.io.

Because DP2 is drawn from a commissioning and early-operations data rather than from the more uniform and systematic LSST, its images and derived products reflect a system that was still being commissioned and tuned. LSSTCam's greater scale and complexity relative to LSSTComCam introduced systematics with no direct counterpart in DP1. For much of the campaign the Active Optics System (AOS) closed loop converged on a slightly mis-calibrated focus target, imprinting a spatially varying defocus across the focal plane that was only corrected on sky in December 2025 and that remains visible in the point-spread function (PSF) of a large fraction of DP2 exposures. Delivered image quality and PSF ellipticity did not consistently meet LSST design requirements in wide-area survey-mode observations by the close of commissioning; and stray- and scattered-light artifacts, sky-background systematics in crowded fields, and a chromatic PSF dependence[26] were identified and only partially mitigated in DP2 processing. These effects, their diagnosis, and the mitigations applied are described in §2 and §5.

Unlike the systematic, rolling revisit pattern that will characterize the LSST, DP2's observing cadence is opportunistic: fields and filters were prioritized to serve AOS and Science Pipelines commissioning rather than to build a representative time series. The Science Validation survey footprint was repeatedly re-scoped over the course of the science validation campaign due to the competing demands of engineering tests and filter availability, which was constrained by hardware issues for part of the campaign. Visit counts and time sampling therefore vary substantially and idiosyncratically from field to field, and DP2 should not be read as a preview of LSST's eventual survey cadence, revisit statistics, or depth-accumulation pattern.

DP2 is being delivered in two phases across the Jul–Dec 2026 window. The first phase, Early DP2 (EDP2), comprises the deep coadd images and all associated catalogs, including new cell-based coadds and the first releases of the ShearObject (§4.3.4, §3.2.8) shape-measurement catalog and the Isolated Star Stellar Motions catalog (§3.2.7). EDP2 is targeted for 2026-07-27. The second phase, DP2, adds the visit-level single-epoch and difference images, completing the full release, and is expected in the Oct–Dec 2026 range. Data products delivered as EDP2 will be incorporated into DP2 unchanged, without additional reprocessing (Guy et al. 2026). This phased delivery follows directly from Rubin's hybrid cloud architecture (O'Mullane et al. 2024a), in which science data products are stored at the United States Data Facility (USDF) and served to users via the US Data Access Center (DAC) in Google Cloud. Deep coadd images and catalogs are hosted directly in Google Cloud and are therefore unaffected by hybrid-cloud readiness, allowing them to ship as EDP2. Visit-level images, however, must be streamed from the USDF, and the necessary capacity is not yet fully in place, introducing timing uncertainty for the second phase. Delivery of the Prompt Data Products follows the same logic. The Prompt Products DataBase (PPDB), Rubin's real-time catalog of transient, variable, and moving-object detections, does not depend on USDF–Google Cloud bandwidth and is therefore prioritized as the second Early-Operations delivery milestone, following EDP2 and preceding the full DP2 release, expected in the Sep–Oct 2026 range. Prompt images (raw images, Processed Visit Images, and difference images), which do require that bandwidth, have been deprioritized relative to EDP2 and the PPDB and are expected to become available only in the Oct–Dec 2026 range, following both.

DP2, like all Rubin Data Previews and Data Releases, is subject to a 2 yr proprietary period; during this time access is restricted to LSST data rights holders (Blum & the Rubin Operations Team 2020) via the RSP (Jurić et al. 2019), Rubin's cloud-based data science platform (Jurić et al. 2019). After the proprietary period, DP2 will be made public, though access through Rubin DACs in the US and Chile will remain limited to data rights holders (Blum & the Rubin Operations Team 2020). Alternative public access mechanisms have not yet been finalized.

The structure of this paper is as follows. Section 2 describes the observatory system and its status at the close of commissioning, the fields included in DP2, and the observing strategy employed. Section 3 summarizes the data products in the release. The processing pipelines are described in Section 4, followed by data validation and performance characterization in Section 5. Section 6 describes the RSP. Section 7 outlines Rubin's community support model. A summary and outlook toward the start of formal LSST Data Releases is given in Section 8.

[26] LSSTCam has no atmospheric dispersion corrector

All magnitudes are quoted in the AB system (Oke & Gunn 1983) unless otherwise specified.

## 2. ON-SKY COMMISSIONING CAMPAIGN WITH LSSTCAM

The LSSTCam on-sky commissioning campaign began on 2025-04-16, following the achievement of LSSTCam first photon on sky, and concluded on 2025-09-21. Its primary objective was to commission and validate and characterize the fully integrated observatory system of the Simonyi Survey Telescope, including the AOS, LSSTCam, and the automated Data Management Subsystem (DMS), all operating together as a single integrated observatory at scale. This was done with a series of focused engineering tests and the execution of a set of Science Validation (SV) surveys. Following a four-week engineering and operations downtime after the conclusion of the SV surveys on 2025-09-21 the observatory was formally handed over from Construction to Operations on 2025-10-25, after which the Science Validation surveys resumed during the Early Operations period. The full DP2 dataset spans both the commissioning SV campaign and this early operations phase and includes observations acquired between 2025-04-16 and the DP2 cut off date of 2026-01-06. The Early Operations period continued until the start of the LSST on 2026-06-29.

### 2.1. *Simonyi Survey Telescope*

The Simonyi Survey Telescope (Stalder et al. 2024) features a compact three-mirror design incorporating an 8.4-meter primary–tertiary mirror (Single piece of glass for Primary Mirror/Tertiary Mirror (M1M3)) fabricated from a single borosilicate monolith, and a 3.5 m secondary mirror (M2). This compact configuration supports a wide 3.5°field of view while enabling exceptional pointing stability, with slew-and-settle times under 5 seconds.

To achieve the scientific goals of the 10-yr LSST, the observatory must maintain high image quality across the full field of view (Ivezić et al. 2019b). This is accomplished through the AOS (Xin et al. 2015; Megias Homar et al. 2024), which corrects, between successive exposures, wavefront distortions caused by optical misalignments and mirror surface deformations, primarily driven by gravitational and thermal loads.

The AOS comprises open-loop and closed-loop components. The open-loop component corrects for predictable distortions using look-up tables as a function of telescope elevation and camera rotator angle. The closed-loop component addresses unpredictable or slowly varying aberrations using curvature wavefront sensing (Thomas et al. 2023) on the eight corner wavefront sensors of the LSSTCam focal plane, which acquire pairs of intra- and extra-focal "donut" images to measure wavefront errors. Corrections are applied to the M2 and camera hexapods and to the shape of M1M3 via its 156-actuator support system, driving the system towards the target optical state within two visits of each measurement. This represents a significant operational advance over the LSSTComCam campaign, during which dedicated wavefront sensors were unavailable and wavefront errors were instead estimated by defocusing the telescope $\pm 1.5$ mm on either side of focus (NSF-DOE Vera C. Rubin Observatory 2025).

During the LSSTCam commissioning campaign, AOS operations were interspersed with dedicated engineering periods focused on wavefront sensing and optical state characterisation, enabling iterative improvements across all components of the system throughout the campaign (Bechtol 2025). The bulk of the DP2 observations were acquired during the commissioning-era SV surveys, a period during which the AOS was undergoing intensive characterisation and iterative improvement. During this time the system contribution to the delivered Point Spread Function (PSF) Full Width at Half-Maximum (FWHM) still exceeded the design specification (Claver & The LSST Systems Engineering Integrated Project Team 2017; NSF-DOE Vera C. Rubin Observatory 2026a,b). As a consequence, the majority of DP2 observations were taken with a larger system contribution to the PSF size budget than the system is capable of delivering under comparable atmospheric conditions, producing larger PSF FWHMs and shallower point-source depths. The residual optical aberrations also manifest as spatially coherent structure in PSF shape residuals across the focal plane, with implications for weak lensing analyses and PSF model validation. The quantitative impact on delivered image quality is characterised in §2.6, and the AOS improvement program and current performance are described in NSF-DOE Vera C. Rubin Observatory (2026a) and NSF-DOE Vera C. Rubin Observatory (2026b).

### 2.2. *LSSTCam*

LSSTCam (Corbett et al. 2024; SLAC National Accelerator Laboratory & NSF-DOE Vera C. Rubin Observatory 2025) is a 3.2-gigapixel mosaic camera with a 9.6 deg$^2$ instantaneous field of view. Its focal plane comprises 25 modular rafts arranged in a 5×5 grid: 21 science rafts, each containing nine 4K×4K Charge-Coupled Device (CCD) sensors, and 4 corner rafts each housing two wavefront sensors and two guider sensors, for a total of 205 sensors: 189 science, 8 wavefront, and 8 guider.

LSSTCam employs sensors from two vendors: Teledyne-e2v semiconductor company (E2V) sensors populate the central science rafts, and Imaging Technology Laboratory (UA) (ITL) sensors populate the peripheral science rafts. Each sensor is subdivided into 16 amplifiers, each reading 512×2048 pixels in parallel. Figure 1 shows the LSSTCam focal plane layout, viewed from above looking down through the camera lenses, including the placement and numbering scheme of sensors ("S") and amplifiers ("C"); a full description is given in Plazas Malagón et al. (2025); Plazas Malagón et al. (2026).

During the commissioning campaign, only 188 of the 189 science sensors were operational; sensor R30/S12 was non-operational throughout and is excluded from all DP2 data products. The eight heavily vignetted corner science sensors (detectors), R01/S00 (0), R03/S02 (20), R10/S00 (27), R14/S02 (65), R30/S20 (123), R34/S22 (161), R41/S20 (168), and R43/S22 (188), are also excluded from all DP2 processing, even though they were functional. Science sensors (detectors) R03/S20 (24), R30/S10 (120), R30/S11 (121), and R42/S00 (172) are additionally excluded from data acquired before 2025-07-03 as the calibration data required to process them were not yet available. Any additional exclusions that were applied during the processing stage are described in §4.

The median amplifier loses approximately 0.45% of its area to static defects, with the mean across amplifiers closer to 0.7% (Bechtol 2025). Overall, more than 99% of the active focal-plane area within the calibrated field is usable (Bechtol 2025). Optical vignetting limits the fully calibrated field radius to 317 mm; a partially vignetted zone extends to 350 mm, beyond which pixels are excluded from science products.

No new static defect classes were identified on either ITL or E2V sensors during commissioning beyond those already identified during commissioning with LSSTComCam (Vera C. Rubin Observatory Team 2026) or in the laboratory (LSST Camera Team et al. 2025; Bechtol 2025). A subset of ITL sensors exhibits "vampire" pixel anomalies and locally elevated readout noise consistent with LSSTComCam commissioning. Dynamic masking, including treatment of E2V edge bleeds and occasional bright-star flare artifacts at extreme detector edges, is ongoing (Bechtol 2025).

### 2.2.1. *Filter Complement*

LSSTCam supports imaging with the same six broadband *ugrizy* filters, spanning 320–1050 nm, as LSSTComCam, using an identical filter design. Its filter exchanger can hold five filters at a time, as opposed to only three for LSSTComCam.

Unlike the LSSTComCam bandpasses, which were built entirely from fiducial mirror, lens, filter, and detector quantum-efficiency curves in the absence of any device-specific measurements, the DP2 standard bandpasses incorporate real measurements where available. The fiducial mirror and lens throughputs are still used, since detector- and mirror-level measurements have not yet been incorporated into DP2. The filter throughput, however, is now the focal-plane average of vendor filter scans, using the same underlying filter design as LSSTComCam but capturing real variation across the field of view. The detector quantum efficiency is likewise the focal-plane average over all the ITL and E2V detectors combined.

ITL and E2V sensors are not identical: ITL devices have somewhat higher quantum efficiency in the red, so the focal-plane-averaged LSSTCam response is marginally lower than the all-ITL response used for LSSTComCam. This chromatic difference is most apparent in the $g$ band, visible even by eye as a two-tone effect across the focal plane. $u$-band throughput is dominated by the atmosphere and filter rather than the detector, and the redder bands are comparatively insensitive to the vendor difference. In absolute terms the ITL/E2V difference is small, much smaller than, for example, the difference between LSSTCam and other instruments such as Dark Energy Camera (DECam), since the two vendors' filters are essentially the same design. It is nonetheless large enough to matter for millimagnitude-level photometric calibration.

Individual per-detector quantum-efficiency curves have not yet been released. The standard passband distributed with DP2 is therefore the focal-plane average, consistent with how the survey is calibrated and appropriate for use with the deep coadds. The full-system throughput of the six LSSTCam filters is shown in Figure 2 and the corresponding transmission curves are provided as a DP2 data product(§3.4.2).

### 2.2.2. *Stray Light*

The compact, 3-mirror design of the Simonyi Survey Telescope (§2.1) is especially sensitive to stray and scattered light[27]. The most pronounced and extended artifacts appear in LSSTCam images when the telescope's boresight is near bright objects such as the Moon, bright stars, planets, or unintended light sources on the sky. Stray light contaminates all astronomical imaging to

[27] Light that does not follow the nominal optical path within the telescope and instrument, originating from scattering off optical and non-optical surfaces or from ghost paths in refractive components such as lenses and filters

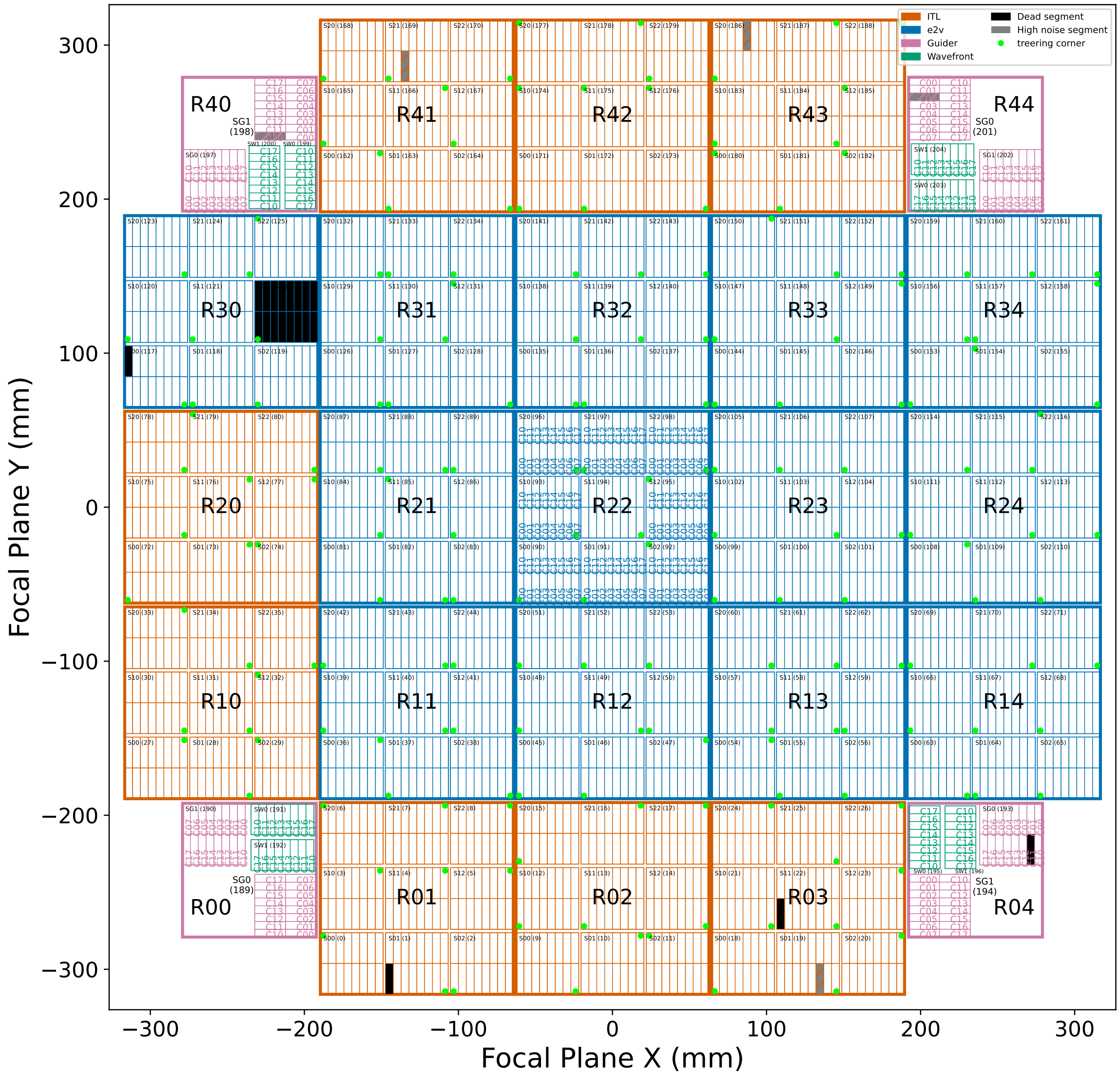


**Figure 1.** LSSTCam focal plane layout, showing the placement and numbering scheme of sensors ("S") and amplifiers ("C"). The view is from above, looking through the LSSTCam lenses onto the focal plane. Each sensor contains 16 amplifiers. Each raft is composed of a 3×3 array of sensors. The detector number for each sensor is indicated in parentheses.

some level, but LSST's low-surface-brightness science is especially sensitive to it.

A significant fraction of these artifacts traced back to the Observatory dome, which lacked its Light-Wind Screen (LWS) to restrict the acceptance angle along the telescopes's elevation axis; many artifacts are expected to resolve once the LWS is comissionined, whereas others will require additional baffling. Some artifacts detected were expected based on an initial stray light analysis in FRED® (Ellis 2009), which correctly predicted direct paths from the primary (M1) and tertiary (M3) mirrors that bypass the secondary mirror and hit the LSSTCam optics and focal plane directly. Other artifacts were new and unexpected, requiring visual inspection, ray tracing simulation, and on-sky tests to confirm their origins (Rodeghiero et al. 2026).

An initial census of stray light artifacts observed in LSSTCam images during initial commissioning (April 2025–May 2026) is reproduced in Table 1 from (Rodeghiero et al. 2026). About 70% of features come from identified, confirmed sources, and 80% already have a possible mitigation strategy. Occurrence rates are based on visual classification of features visible by eye; for some, fractional area is a rough estimate due to variation across frames.

The most frequent and impactful feature, the *scratched tape* feature (Drlica-Wagner et al. 2026), originates from the M1 direct path that skips M2 and M3 reflections, and was visible in ∼7.4% of image frames

**Table 1.** Occurrence rate of stray and scattered light artifacts identifiable by visual inspection of $\sim 116{,}000$ full focal plane mosaic images. Rates are not expected to be complete nor pure; impacted area uses average feature size. The most prevalent artifact, the *scratched tape*, was effectively mitigated through installation of the mid-level baffle extension (Drlica-Wagner et al. 2026). * Fractional area of the artifact is variable; † artifact origin to be confirmed; ϒ preliminary estimate.

| Artifact | Source | Mitigation | Occurrence Rate | Fractional Area |
|---|---|---|---|---|
| Scratched Tape | M1 direct path | Mid-Level Baffle Extension | 7.39% | 2.2% * |
| Smudge | M1 direct path | Mid-Level Baffle Extension | 0.19% | 0.28% |
| Muddy Shoe | M1 direct path | Mid-Level Baffle Extension | 0.15% | 2%* |
| Pillow | Unknown | LWS installation | 0.26% | 9%* |
| Nibbled Disc | Unknown | Mid-Level Baffle Extension | 0.10% | 1.06 - 10%* |
| L2 Bananas | L2 support pads | LWS installation | 0.73% | 0.7% |
| Rivet Holes | Holes in Top-end Baffle | Sealing of holes | 0.45% | – |
| Sharp Arcs | L2 edges † | TBD | TBD | 1% ϒ |
| Submarine | L1/L2 ghosts † | LWS installation (TBC) | 0.7% ϒ | 8.5%* |
| L3 Horseshoe | L3 chamfer & mount | Additional L3 edge baffle | 0.09% | 8.8% |
| All artifacts | — | — | 10.00% | — |
| Post Mid-Level Baffle Extension installation | | — | 2.6% | — |



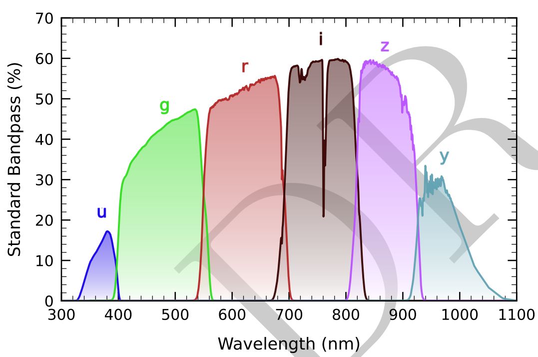


**Figure 2.** LSSTCam standard bandpasses, illustrating full-shown stem throughput. The bandpasses include a standard atmosphere at an airmass of 1.2, telescope optics, camera surfaces, and mean ITL and E2V detector quantum efficiency. The corresponding transmission curves are provided as a DP2 data product.

with an off-axis range of visibility between $\sim 19\overset{\circ}{.}8$ and 22°. An annular extension of the telescope's mid-level baffle ring, installed between February and April 2026, elimintated this feature from LSSTCam images (Drlica-Wagner et al. 2026). Other features such as the *L3 Horseshoe*, arise from edge diffraction at LSSTCam's tertiary lens (L3) near the field-of-view edge ($\sim 1\overset{\circ}{.}99$ and $2\overset{\circ}{.}23$). Ray tracing simulations also revealed scattered light from the Secondary Mirror (M2) optical baffle, whose residual reflectance in the longer wavelength bands ($i$, $z$, $y$) can reach up to 30%.

Despite the preventative measures taken, $\sim 10\%$ of early image frames suffer from stray light contamination, dropping to $\sim 2.6\%$ following the mid-level baffle extension installation (Drlica-Wagner et al. 2026). The most severe issues were structured artifacts on the focal plane, largely stemming from the incomplete status of the LWS. Rubin is implementing new physical baffling and re-painting to address the remaining artifacts (Pollek et al. 2026; Taranto et al. 2026).

### 2.3. *Calibration System*

Unlike the LSSTComCam campaign, during which the Rubin calibration system had not yet been installed and flat fielding relied entirely on twilight flats (Vera C. Rubin Observatory Team 2026), the LSSTCam commissioning campaign had access to the Rubin flat field illumination system (Fagrelius & Rykoff 2025). The baseline calibration plan for the as-built system is described in Fagrelius & Rykoff (2025); a full description of Instrument Signal Removal (ISR) procedures and calibration data products is given in Plazas Malagón et al. (2025).

The calibration system comprises a dome flat field screen, a Collimated Beam Projector (CBP), a National Institute of Standards and Technology (USA) (NIST)-traceable silicon photodiode electrometer, and the AuxTel for atmospheric monitoring. During the commissioning campaign, the CBP was deployed at scale on LSSTCam for the first time, providing in-situ measurements of the transmission profiles of five of the six LSST broadband filters, *ugrizy*, across multiple primary mirror pupil positions, and characterising the angle-of-incidence-dependent blue-shift of the filter bandpass edges across the focal plane (Amouroux et al. 2026). Achieving absolute throughput calibration at the milli-magnitude level required by the LSST requires a more precise characterisation of the CBP transmission and a partial-pupil vignetting model; both are the subject of ongoing work (Amouroux et al. 2026).

Calibration products generated for DP2 processing include bias frames, dark frames, dome flat fields, Photon Transfer Curves (PTCs), cross talk matrices, brighter-fatter kernels, linearity corrections, and Charge Transfer Inefficiency (CTI) models (Plazas Malagón et al. 2025). Dome flat fields were used as the primary flat field source, supplemented by a star-flat-based illumination correction and, where dome illumination was unavailable, twilight flats (Plazas Malagón et al. 2025). Certified calibration products are versioned in the USDF archive and transferred to the summit as needed. The full suite of calibration and throughput data products included in DP2 is described in §3.

### 2.4. *Science Program Observations*

The DP2 dataset spans three distinct phases of the commissioning and early operations program: small-field observations acquired shortly after the LSSTCam First Photon milestone in support of the Rubin First Look event; the commissioning-era SV surveys; and pre-LSST observations acquired during the Early Operations period between the construction-to-operations handover on 2025-10-25 and 2026-01-06. The sky distribution of DP2 visits across all three phases is shown in Figure 3.

All observations were taken using the Feature-Based Scheduler (FBS), which scheduled visits adaptively to balance sky coverage, filter distribution, airmass, and lunar phase constraints (Jones 2026). Data collected between the end of the DP2 observing window on 2026-01-06 and the formal start of the LSST on 2026-06-29 will be processed and made available to LSST data rights holders through the same RSP services. The format and timing of that release are still to be determined; planning will begin once DP2 has been delivered (Guy et al. 2026)

#### 2.4.1. *Small-Field Observations*

Several small fields were observed prior to the start of the SV surveys, with some observations overlapping the early portion of the SV survey period. All small-field observations, including those from the First Look campaign, are included in DP2. Observations were typically acquired in sequences of at least 10 visits per bandpass, cycling through the available filters, with field selection guided by commissioning priorities and community input. Several of these fields, including Cosmic Evolution Survey (COSMOS), M49, the Prawn Nebula region, and Rubin_SV_225_-40, received substantial numbers of visits and are among the most deeply observed regions in DP2. The COSMOS field was first observed as part of the small-field campaign but was not accessible during the main SV survey period; additional observations were later acquired during the Early Operations phase.

A summary of the data collected during the small-field surveys, including visit counts and median delivered image quality expressed as the PSF FWHM, is provided in Table 2

**Table 2.** Number of visits per small-field science verification region and per band in the DP2 dataset. The All column gives the total visits across all bands; the All row gives the total visits across all fields.

| **Field** | **Band** | | | | | | **All** |
|---|---|---|---|---|---|---|---|
| | *u* | *g* | *r* | *i* | *z* | *y* | |
| Abell_2764 | 0 | 6 | 0 | 2 | 0 | 0 | 8 |
| DESI_SV3_R1 | 0 | 0 | 0 | 3 | 0 | 0 | 3 |
| M49 | 234 | 280 | 378 | 213 | 0 | 0 | 1105 |
| Prawn | 193 | 159 | 139 | 91 | 30 | 0 | 612 |
| Trifid-Lagoon | 231 | 201 | 129 | 119 | 9 | 6 | 695 |
| New_Horizons | 37 | 53 | 77 | 116 | 80 | 33 | 396 |
| Rubin_SV_212_-7 | 0 | 139 | 187 | 123 | 0 | 0 | 449 |
| Rubin_SV_216_-17 | 0 | 27 | 28 | 59 | 0 | 0 | 114 |
| Rubin_SV_225_-40 | 307 | 524 | 391 | 381 | 239 | 104 | 1946 |
| Rubin_SV_280_-48 | 30 | 29 | 26 | 30 | 29 | 0 | 144 |
| Rubin_SV_300_-41 | 0 | 8 | 0 | 0 | 30 | 0 | 38 |
| Rubin_SV_320_-15 | 12 | 47 | 104 | 259 | 211 | 77 | 710 |
| **All** | 1044 | 1473 | 1459 | 1396 | 628 | 220 | 6220 |

#### 2.4.2. *Science Validation Surveys*

Plots and estimates of coadded depth still need to be added

The SV surveys began on 2025-04-16 and comprised two interleaved components within a single FBS configuration, a wide survey and a deep survey.

The Wide Survey covered an initial footprint of $\sim 3{,}000\,\mathrm{deg}^2$ tracing the ecliptic plane from dense regions of the Galactic Bulge through low-dust regions of the planned LSST Wide-Fast-Deep (WFD) footprint. Approximately halfway through the SV period, the footprint was consolidated to $\sim 750\,\mathrm{deg}^2$, and further to $\sim 300\,\mathrm{deg}^2$ in the final week of commissioning to concentrate visits and achieve the depth required for difference-image testing. As a result, median visit counts and coad-

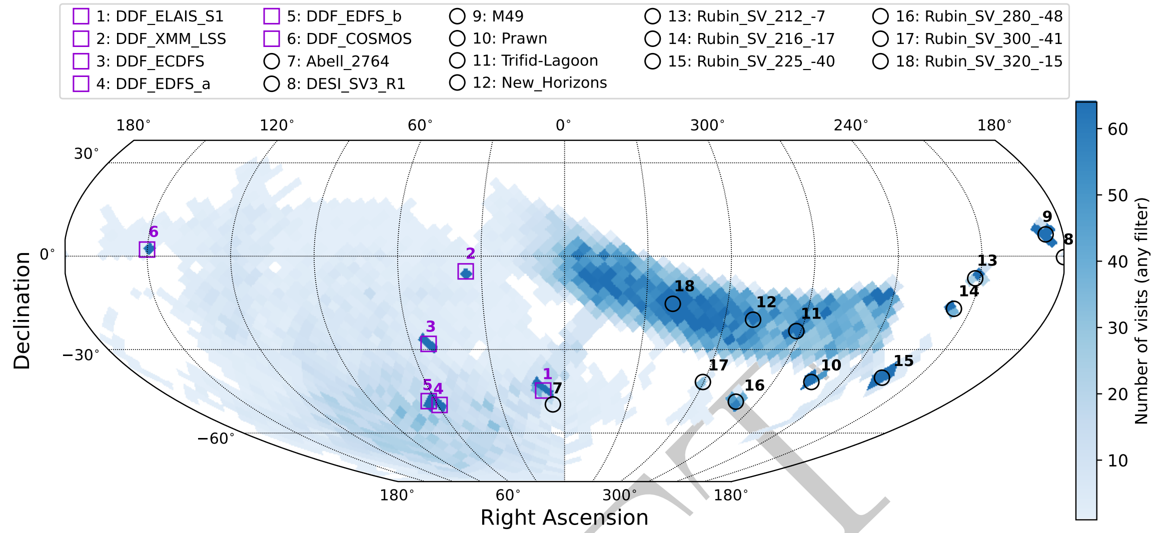


**Figure 3.** Sky distribution of DP2 visits across the three components of the observing program. Small-field survey areas are marked with black circles and Deep Drilling Fields (DDFs) with violet squares. The SV wide survey covers $\sim 3,000\,\mathrm{deg}^2$ tracing the ecliptic plane, with visit density increasing toward the $\sim 300\,\mathrm{deg}^2$ core. Pre-LSST Early Operations observations appear as sparser, predominantly $i$-band coverage extending across a larger area of the LSST footprint.

ded depth increase substantially toward the inner regions of the footprint: the full 3,000 $\mathrm{deg}^2$ region reaches a median of 38 visits across all bands (median $5\sigma$ coadded point-source detection limit; hereafter $m_5 \approx 25.0$), the 750 $\mathrm{deg}^2$ subset reaches 56 visits ($m_5 \approx 25.8$), and the 300 $\mathrm{deg}^2$ core reaches 64 visits ($m_5 \approx 25.7$).

Tables of n visits and depth in WFD area still to come

The Deep Survey targeted four of the five planned LSST DDFs:Extended Chandra Deep Field-South Survey (ECDFS), Euclid Deep Field South (EDFS) (fields A and B), ELAISS1, and COSMOS. The FBS interleaved DDF observations with the wide survey, scheduling visits to each field according to the prescribed DDF observing cadence. ELAISS1 received the largest number of visits, 539 across all six *ugrizy* bands, and is the only DDF observed in all bands during the SV survey. COSMOS was not accessible during the SV period as described above. The remaining DDFs were observed primarily in *griz*, as the $u$ and $y$ filters were unavailable after 2025-08-10 due to a hardware issue with the filter exchanger. The dither pattern for the DDFs evolved during the campaign; early observations used only small translational dithers with no intra-night rotation, while later observations introduced both translational and rotational intra-night dithers. Table 3 provides the median numbers of visits within each of the subsets of the wide SV area.

**Table 3.** Number of visits per Deep Drilling Field and per band in the DP2 dataset. The All column gives the total visits across all bands; the All row gives the total visits across all DDFs.

| **Field** | **Band** | | | | | | **All** |
|---|---|---|---|---|---|---|---|
| | $u$ | $g$ | $r$ | $i$ | $z$ | $y$ | |
| DDF_ELAIS_S1 | 57 | 153 | 128 | 232 | 141 | 59 | 770 |
| DDF_XMM_LSS | 42 | 15 | 8 | 47 | 10 | 10 | 132 |
| DDF_ECDFS | 10 | 61 | 50 | 96 | 61 | 11 | 289 |
| DDF_EDFS_a | 7 | 31 | 33 | 59 | 33 | 7 | 170 |
| DDF_EDFS_b | 5 | 37 | 33 | 57 | 33 | 7 | 172 |
| DDF_COSMOS | 105 | 87 | 173 | 162 | 110 | 77 | 714 |
| **All** | 226 | 384 | 425 | 653 | 388 | 171 | 2247 |

### 2.4.3. *Early Operations Observations*

Following the transition to Operations on 2025-10-25, FBS-driven observations resumed during the Early Operations optimization period. Observations during this period, which ran from 2025-10-25 through 2026-01-06, were a mixture of engineering activities and FBS-driven survey observations aimed at improving image quality and survey speed in preparation to start the LSST, and building difference-imaging templates across the full LSST footprint.

The pre-LSST observations are characterised by a lower and less regular cadence than the SV surveys, reflecting frequent engineering interruptions. The FBS was configured with a single-visit template tier as a scheduling priority, and on any given night would typically concentrate observations in a single band rather than interleaving filters, to maximise sky coverage for template building. The *i* band was preferred due to its high throughput and suitability for difference imaging. A concentrated block of *i*-band observations in late October 2025 reflected deliberate AOS engineering tests that held the filter fixed to isolate optical state changes from filter-induced perturbations, producing the prominent *i*-band coverage regions visible in Figure 3. Pre-LSST visits not overlapping existing SV survey regions were not combined into coadds but are included in the DP2 visit dataset.

### 2.5. *Data Processing*

The Rubin DMS is responsible for the acquisition, transfer, archiving, processing, and distribution of all data produced by the observatory. It operates as a fully automated end-to-end system, ingesting raw pixels from the LSSTCam focal plane within seconds of shutter close and delivering science-ready data products to a global community of researchers within 120 secs.

The data products in DP2 are the result of Data Release Processing (DRP) (§4), which reprocesses the full accumulated image archive to produce the deep, well-calibrated data products that form the basis of each Data Release, and are delivered to the community via the RSP (§6). During the LSSTCam commissioning campaign, the DMS was operated at full scale for the first time, transferring, archiving, and processing all LSSTCam images and validating the end-to-end data flow from summit to RSP.

#### 2.5.1. *Data Acquisition, Transfer and Archive*

Following each exposure, the Data Acquisition System (DAQ) system at the summit stores the raw image data as a set of 205 files per exposure: 189 science sensors[28], 8 wavefront sensors, and 8 guiding sensors on the LSSTCam focal plane. Uncompressed, a full science exposure totals $\sim 8.2$ GB; after compression, $\sim 4$ GB.

At the summit, each exposure is processed in near real time by the Rapid Analysis pipeline, performing ISR (§4.1.1) and source detection (§4.1) on a 30–60 second timescale to provide quick-look image diagnostics and support real-time AOS assessment. The resulting per-visit quality metrics formed the basis of the visit quality characterisations used to define the DP2 coadd inputs. Summary visit and detector-level metadata are available to users as DP2 data products (§3.5) and are also available in an opsim-format SQLite database ()§6.5.2).

#### 2.5.2. *LSST Science Pipelines*

The DP2 science-ready data products in are the result of dedicated DRP campaigns executed at the USDF at SLAC National Accelerator Laboratory (SLAC). DP2 processing was done using version v30.0.8 of the LSST Science Pipelines (Rubin Observatory Science Pipelines Developers et al. 2026) and follows a four-stage structure. Stage 1 performs single-visit processing, including ISR, source detection, astrometric and photometric calibration, and PSF characterisation. Stage 2 performs global recalibration, Gaussian-process modelling of atmospheric turbulence, proper motion and parallax fitting, and PSF re-estimation using a standard star set. Stage 3 performs coaddition and measurement on coadded images, including the generation of the Object catalog. Stage 4 performs Difference Image Analysis (DIA) and solar system object processing, and the generation of light curves via forced photometry. Each stage underwent continuous data quality monitoring and a formal verification and validation review before processing advanced to the subsequent stage. A full description of the processing pipelines, algorithms, and configuration used for DP2 is given in §4.

### 2.6. *Delivered Image Quality*

The delivered image quality at Rubin Observatory reflects contributions from the atmosphere, dome environment, telescope optics, AOS, and LSSTCam. The design specification for the median delivered PSF FWHM, referenced to zenith at 550,nm, is $0''.7$ (Ivezić & The LSST Science Collaboration 2018), corresponding to the measured median atmospheric seeing of $0''.72$ at Cerro Pach'on (NSF-DOE Vera C. Rubin Observatory 2026b) added in quadrature with a total system contribution of $0''.35$ (NSF-DOE Vera C. Rubin Observatory 2026b). During the period covered by the DP2 dataset, the system had not yet consistently met this design specification (NSF-DOE Vera C. Rubin Observatory 2026b; Bechtol 2025). The primary excess contributions are from the optical system, including residual AOS errors and mirror thermal gradients, as well as from dome seeing. The current total system contribution to PSF FWHM is approximately $0''.55$–$0''.60$, compared to the design allocation of $0''.40$. At the median survey airmass of 1.2, an atmospheric contribution of $0''.72$ at zenith corresponds to $\sim 0''.80$; combined in quadrature with the current system contribution, this is consistent with the observed median delivered image quality of $\sim 1''.0$ PSF FWHM in the *riz* bands during the DP2 observations (NSF-DOE Vera C. Rubin Observatory 2026a).

[28] Only 188 science sensor files were generated per exposure as sensor R30/S12 was non-operational throughout; see §2.2.

Image quality tables and plots still to come

## 3. OVERVIEW OF THE CONTENTS OF RUBIN DP2

In this section we describe the data products that are included in DP2. Data rights holders can access all of these data products via one or more of the data services outlined in §6. For programmatic reasons, the precise names of the data product can differ slightly depending on which data service is being used to retrieve them. For this reason, we use descriptive names throughout this section and provide their equivalent names used in the data services in Table 4. We note that some data products are not available in all data services; such cases are indicated in Table 4.

As with DP1, all the data products described in this section fall under one of five categories:

- Science Images, including pre-processed Raw Images, Visit Images, and Deep Coadd images;
- Catalogs of astrophysical Sources and Objects detected and measured in the aforementioned science images. We also provide the reference catalog that was constructed to astrometrically and photometrically calibrate the data during processing;
- Maps, which provide non-science-grade visualisations of the data;
- Ancillary Data Products, such as the values of the configuration parameters that were used during data processing, processing logs, and calibration data products;
- Metadata, in the form of tables containing information about each visit and visit image, such as visit pointing coordinates, exposure times, and data quality summary statistics.

We expect Science Images and Catalogs to be the most commonly used categories for science, while Maps, Ancillary Data Products, and Metadata remain important for facilitating science exploitation via visual inspection, as well as for transparency, Education and Public Outreach (EPO), media, and citizen-science activities such as Zooniverse. While the categories of data products in DP2 are the same as in DP1, this second Data Preview contains a wider range of data product types compared to DP1. This is principally due to the science processing pipeline being in a more mature state compared to the first data preview. In the following subsections we clearly indicate those data product types which are new to DP2.

As with DP1, the science processing pipelines use a skymap to structure how processing is organized spatially. The skymap divides the sky into a series of equal-area tracts, which are then subdivided into a number of patches. The tracts are arranged in rings of constant decl., starting at the southern pole and having an index number that increases first by increasing R.A. and then by increasing decl.. The skymap that was generated for DP2 processing contains a total of 18938 tracts spanning the whole sky, 2174 of which were covered during the science validation observing campaign, including the small field surveys. Each DP2 tract covers approximately 2.8 $\deg^2$, and contains 10 $\times$10 patches, both the same as for the DP1 skymap. Tracts are defined to have equal area, and so the amount by which neighboring tracts overlap varies with decl., with tracts near the celestial poles overlapping more than those near the celestial equator. The minimum area of overlap between tracts is $1'.0$. Within a tract, the area of overlap between neighboring patches is a constant $60''.0$. As a result, each patch covers an area of 0.034 $\deg^2$, which is larger than the tract area divided by the number of patches in a tract. Excluding the overlap region, the area covered uniquely by a single patch is 0.028 $\deg^2$.

Although the tract and patch configuration is unchanged between DP1 and DP2, the two skymaps differ at the cell level for Deep Coadds. Patches are further subdivided into cells for the sole purpose of enabling the construction of cell-based coadds with a continuous PSF (§3.1.3). Unlike patches, cells are not exposed to the user, and one cannot specify nor retrieve an individual "cell" from an image. Instead, cells are used during the construction of what are termed cell-based coadds. A region of an input image is included in a given cell if, and only if, it covers the entire cell. This guarantees that every pixel within a cell is built from exactly the same set of input images, so the PSF is well-defined and effectively constant across the cell, with discontinuities confined to the cell boundaries. Such cell-based coaddition was not implemented during the production of DP1, and was only used to produce the Deep Coadds in DP2. Template Coadds in DP2 do not utilize cells. In the case of DP2 Deep Coadds, there are 22 $\times$22 cells per patch, including the patch overlap region (the patch overlap region is 1 cell wide). There is no overlap between neighboring cells.

### 3.1. *Science Images*

Science images are images taken with the intent of scientific analysis of astronomical objects, as opposed to calibration images (e.g., bias, darks, flats) or engineering images, which serve to calibrate or test the instrument

**Table 4.** Summary of dataset types released in EDP2, together with the names by which they are referred to within the QServ database and the Data Butler.

| Dataset Type | QServe Name | Butler Name |
|---|---|---|
| Deep Coadd | – | deep_coadd |
| Object | dp2.Object | object |
| Isolated Star Stellar Motions | dp2.IsolatedStarStellarMotions | isolated_star_stellar_motions |
| Object Shear | dp2.ShearObject | object_shear_all |
| Source | dp2.Source | source |
| Forced Source | dp2.ForcedSource | object_forced_source |
| Forced Source at DIA Object | dp2.ForcedSourceOnDiaObject | dia_object_forced_source |
| DIA Object | dp2.DiaObject | dia_object |
| DIA Source | dp2.DiaSource | dia_source |
| Solar System Object | dp2.SSObject | ss_object |
| MPC Orbit | dp2.mpc_orbits | – |
| Calibration Reference Catalog | – | the_monster_20250219 |
| Survey Property maps | – | deepCoadd_<PROPERTY>_consolidated_map_<STATISTIC> |
| HiPS Maps | – | – |
| Skymap | – | skyMap |
| Standard Passbands | – | standard_passband |
| Object Scarlet Models | – | object_scarlet_models |
| Deep Coadd Input Summary | – | deep_coadd_input_summary |
| Visit | dp2.Visit | visit_table |
| Visit Detector | dp2.VisitDetector | visit_detector_table |
| Visit Summary | – | visit_summary |



and telescope regardless of whether they happen to be pointed at the night sky. DP2 provides all images in the acquisition and processing chain, from the raw, unprocessed science and calibration images through to fully calibrated visit images, coadds, and template images. Most science analyses will start from the Deep Coadd images and the single-epoch Visit images. While it is possible for users to recreate their own Processed Visit Images (PVIs) from the raw and calibration data, we do not expect anyone will need to do so. One important difference between DP1 and DP2 is the format of the Science Images (Bosch 2026). In the case of DP2, the format of the datasets are such that they are more compatible with external packages, in particular Astropy. It is, however, straightforward to convert back to "legacy" DP1-style datasets for backwards-compatability.

#### 3.1.1. *Raw Images*

Raw Images are those that are produced by the camera following an exposure. They are the initial input science images to the pipelines and have not undergone any processing. Each active detector in LSSTCam (§2.2) produces one Raw Image per exposure. Each exposure therefore results in up to 180 Raw Images, although some exposures taken during the observing campaign resulted in fewer than 180 Raw Images due to the problematic detectors described in §2 and occasional readout failures.

All LSSTCam detectors have 16 amplifier segments, arranged in two rows of eight, each reading out its own $\sim 1$ megapixel region in parallel to enable the fast ($\sim 2$ second) full focal-plane readout (§2.2). These sections are evident in the Raw Images. Despite LSSTCam containing detectors from two different vendors, all Raw Images have the same dimensions: $4608 \times 4096$ pixels. However, the two types of detectors have slightly different numbers of overscan and pre-scan pixels, meaning the number of exposed pixels differs between detector type. This difference is reflected in the dimensions of the Visit Images (§3.1.2), which have had the prescan and overscan regions removed during data processing.

#### 3.1.2. *Visit Images*

**Note:** Visit Images are NOT included in EDP2, but we describe them here as it provides useful context for the dataset types that are included in EDP2. Visit Images will be included in DP2; the number of Visit Images in DP2 will be provided at that time.

Visit Images are the first scientifically useful images to be produced by the science processing pipeline. They are the result of removing instrument signatures from the raw input images, followed by multiple rounds of background subtraction and photometric and astrometric calibration (see §4). There is one Visit Image for every raw input image that successfully passed through the first stage of the science processing pipelines. As such, there are up to 180 Visit Images for each of the exposures taken during the observing campaign.

During ISR, the science processing pipeline removes the overscan regions from the input raw images. The Visit Images therefore only contain the pixels that were exposed during the observation: 4072 ×4000 pixels in the case of the ITL detectors, and 4096 ×4004 pixels in the case of the e2v detectors. Unlike raw images, Visit Images consist of three image planes: a science image plane, a variance plane, and a mask plane. The science image is the processed exposed image, whereas the variance image persists the uncertainty associated with the value of the corresponding pixel in the science image. The mask image indicates pixels that are flagged for various different reasons, such as saturated pixels, pixels associated with detected sources, bad pixels, etc. All three images contain the same number of pixels in each dimension.

### 3.1.3. *Coadds*

Once produced, individual Visit Images covering the same patch of sky (as defined by the skymap) in the same band are coadded to form coadd images. Two different types of coadd images are included in DP2: Deep Coadds and Template Coadds, both of which are described in the following subsections. The production of each type of coadd image follows broadly the same process, which is described in detail in §4.3.2 but, to summarize, it involves the selection of Visit Images (in terms of patch coverage and band, and image quality), the warping of those Visit Images onto a common pixel grid, and the coaddition of those warped images. The main difference that distinguishes each type of coadd are the selection of the Visit Images. As a result, the structure of the datasets is the same for each type: each coadd contains three planes corresponding to the science image plane (i.e., the coadd), the variance plane, and the mask plane, with all three containing the same number of pixels in each dimension. Up to two different background models are included in the coadd datasets: the background that was subtracted to create the final coadd and, in the case of the *gri* band Deep Coadds, a background that can be added-back to create "Visualization Coadds" which, while not recommended for science, can be useful for visualization. Each coadd also contains a position-dependent PSF model, WCS information, plus various other metadata.

The same skymap was used for each type of coadd in DP2, meaning they all cover the same on-sky area, including some overlap with their neighbouring coadds (see the introduction to this section). A single coadd, irrespective of type, contains 3300 ×3300 equal-sized square pixels, including overlap; each pixel is 0″.2 on each side. There are up to six coadds of each type per patch - one for each of the six LSSTCam bands. While the sky coverage of coadds is similar between types and bands, they are not exactly the same due to the different Visit Image selection functions for each type, and differences in sky coverages in each band.

The number of Visit Images that contribute to a coadd varies across the patch. This is true even if a patch is fully covered by all the visits that targeted that area because of the presence of chip gaps between neighbouring detectors. The Survey Property maps (see §3.3.1) can be consulted to gain an insight into the total exposure time, and thus the number of contributing Visit Images, at any point within a coadd. Since each coadd image always covers an entire patch, it is not uncommon for regions of a coadd image to have zero Visit Image coverage. This is especially the case on the edges of an observed field. While such regions may contain pixels with apparently valid values (i.e., non-zero or non-`NaN`s), they will be flagged as `NO_DATA` in the mask image plane. It is therefore crucial that the mask image plane is consulted when analysing any type of coadd.

*Deep Coadds*—

Coadd background description

As the name suggests, Deep Coadds go down to the faintest magnitudes of the three types of coadd. Two basic selection criteria were used to select visits for inclusion in Deep Coadds. The first criterion is seeing-based and selects visits with a mean (over all detectors) PSF FWHM<1″.7. The second criterion is depth-based and uses the concept of "effective exposure time", which is the exposure time that would be needed under nominal conditions to achieve the same signal-to-noise as the visit in question (Bernstein et al. 2016). Our specific selection uses the transparency component of effective exposure time, and selects visits in which the whole-system transparency was such that the transparency component of the effective exposure time was $> 0.75$ times the actual exposure time (corresponding to $\sim 0.16$ mag of extinction relative to nominal conditions). Following the selection of visits, a second round of selection is applied at the detector (i.e., Visit Image) level. These per-detector selection criteria are summarised in Table 5. Together, these selections aim to maximise the depth of the coadd while still ensuring good image quality. As highlighted at the start of this subsection, a key difference between the Deep Coadds released in DP2 versus those released in DP1 is that the DP2 Deep Coadds are cell-based whereas the DP1 Deep Coadds were not.

There are a total of 925460 Deep Coadds in DP2.

*Template Coadds*—

This needs checking. What other criteria are used for selection beyond just PSF? Also need to describe background.

Template Coadds are created, principally, to act as reference images in the process of image subtraction to produce Difference Images (see §??). When performing image subtraction, the point spread function (PSF) of the reference image is degraded to match that of the science image. For this reason, it is beneficial for the reference image to have as small a PSF as possible whilst also ensuring they are deep enough so as to ensure they do not introduce substantial noise during subtraction (i.e., the noise budget is dominated by the single visit image, as opposed to the template coadd). When selecting Visit Images for Template Coadds, only images with a PSF FWHM<1.″4 are considered. Of the Visit Images that meet that PSF FWHM criterion, the third that cover a given region in a given band with the smallest PSF are selected for coaddition. If that third corresponds to fewer than 12 Visit Images (i.e., there are fewer than 36 Visit Images that cover the region in a given band), then the 12 Visit Images with the smallest PSFs are selected. If there are fewer than 12 Visit Images covering a region with a given band, then all Visit Images (with PSF FWHM<1.″4) are selected. This strategy optimizes for image quality in regions that are well-covered by Visit Images, while also ensuring good Template Coadd coverage in regions that are poorly-covered by Visit Images.

In total, there are 0 Template Coadds in DP2.

## 3.2. *Catalogs*

Catalogs contain data for detections in the images described in §3.1, except the `Calibration` catalog, which contains reference data obtained from previous surveys. Observatory-produced metadata tables are described in §3.5. Rubin adopts a naming convention in which the term "object" refers to an astrophysical object, whereas the term "source" refers to a single-epoch detection of that object. As such, sources are associated with detections in single-epoch Visit Images and Difference Images, whereas objects are associated with detections in multiepoch coadds. Photometric quantities in catalogs are stored as fluxes rather than magnitudes, with fluxes measured in nanojanskies (1 nJy = $10^{-35}$ $\mathrm{Wm^{-2}Hz^{-1}}$ ). Fluxes are preferred for multiepoch observations, as they can be averaged across epochs, unlike magnitudes. Additionally, flux measurements on difference images (§3.1) are computed relative to a template and can therefore be negative, particularly for faint sources in the presence of noise.

The catalogs described below each vary in terms of their specific columns but most contain: one or more unique identification numbers, positional information, multiple types of flux measurements (e.g., aperture fluxes, PSF fluxes, Gaussian fluxes, etc.), and a series of Boolean flags indicating characteristics such as saturation or cosmic-ray contamination for each source or object. The solar system catalogs SSObject and SSSource deviate from this general structure in that they instead contain orbital parameters for all known asteroids.

Where applicable, quantities are prefixed with the band in which they were measured, and all measured properties are reported with their associated 1$\sigma$ uncertainties. For example, `g_ra` and `g_raErr` in the Object catalog refer to R.A. and its uncertainty, measured in the $g$ band. Fluxes for various apertures are provided together with an uncertainty and a flag, and named in the format `[band_]ap[size]Flux`, where `[size]` is the aperture diameter in pixels, and `[band_]` is used where applicable (e.g., in the object catalog, but not in the source catalog; see §3.2.1 and §3.2.3). For example, `g_ap03Flux`, `g_ap03FluxErr`, `g_ap03Flux_flag` provide the flux, uncertainty and a flag measured within a 3.0 pixel aperture in the $g$ band. Similarly for flux measurements using difference algorithms, e.g. `g_psfFlux` provides the flux derived using the PSF model as a weight function, forced on $g$ band.

All of the catalog types that were released as part of DP1 have been retained in DP2, although in some cases additional columns have been added for DP2; we highlight prominent additional columns in the subsections below. There are also a number of additional catalogs in DP2 that do not have DP1 equivalents. Throughout this section, we provide high-level overviews of the contents of each catalog, rather than listing and describing each column. Instead, full searchable catalogue schemas, including column descriptions, are provided at https://sdm-schemas.lsst.io/.

All catalogs except the Calibration Reference Catalog are stored in the Qserv database (§6.5.1) and accessible via Table Access Protocol (IVOA standard) (TAP), with an online DP2 catalog schema available

**Table 5.** Per-detector PSF-quality selection cuts for inclusion in Deep Coadd images.

| Parameter | Description | Numerical cut |
|---|---|---|
| PSF FWHM | Same FWHM calculation as the visit-level cut, evaluated per detector rather than averaged over the visit. | $< 1.8$ arcsec |
| Median ellipticity residual | Median of $\sqrt{e_1^2 + e_2^2}$ between the PSF model and measured star shapes.[a] | $< 0.19$ |
| `psfStarScaledDeltaSizeScatter` | Robust scatter of PSF size residuals, scaled by the median PSF size. | $< 0.07$ |
| `psfTraceRadiusDelta` | Max−min of the PSF trace radius evaluated on a grid across the unmasked detector pixels. | $< 4.4$ pixels |
| `psfApFluxDelta` | Max−min of the model PSF aperture flux (normalized) across the unmasked detector. | $< 1.6$ |
| `psfApCorrSigmaScaledDelta` | Max−min of the PSF aperture correction across the detector, scaled by `psfSigma`. | $< 0.15$ |
| `starEMedian` | Median star ellipticity, $\sqrt{e_1^2 + e_2^2}$.[a] | $< 0.2$ (all bands) |
| `starUnNormalizedEMedian` | Same, computed from unnormalized second moments.[b] | $< 2.3$ (all bands) |
| `nPsfStar` | Number of PSF-candidate stars available on the detector; below this the PSF model is not considered well-constrained. | $\geq 6$ per band |

[a] $e_1 = (I_{xx} - I_{yy})/(I_{xx} + I_{yy})$, $e_2 = 2I_{xy}/(I_{xx} + I_{yy})$, where $I_{xx}$, $I_{yy}$, $I_{xy}$ are the adaptive second moments of the PSF/star.

[b] Unnormalized equivalent: $\sqrt{(I_{xx} - I_{yy})^2 + (2I_{xy})^2}$, i.e. without dividing by $I_{xx} + I_{yy}$.

at https://sdm-schemas.lsst.io/dp2.html. Catalog data are also accessible via the Data Butler (see §6.2.2).

#### 3.2.1. *Source*

The Source catalog contains the measured properties of all sources that are, prior to deblending, detected at a $> 5\sigma$ significance in each of the Visit Images (§3.1.2). The detections in the Source catalog have all undergone deblending; in the case of blended sources, only the post-deblended sources are included in the Source catalog. It is therefore possible for the Source catalog to contain entries in which none of the different types of photometric measurements (e.g., aperture, PSF, gaussian) have flux-to-error ratios greater than five.

Each entry in the source catalog has a unique identification number (`sourceId`), positional information, various photometric measurements (aperture, PSF, gaussian), shape and extendedness information, and a series of flags. Because each entry in the source catalog is derived from a detection in a single epoch observation taken with a single filter, each entry also has an associated band.

The Source catalog contains data for 18 billion sources in DP2.

#### 3.2.2. *DIA Source*

The DIA Source catalog contains data on sources detected at $> 5\sigma$ significance in the Difference Images, including those associated with known solar system objects. Following the detection and measurement of each DIA source, an attempt was made to associate it with a DIA object (see §3.2.4). If no suitable DIA object was found, then a new entry was created in the DIA Object catalog. So as well as being indexed on a unique `diaSourceId` identification number, each entry in the DIA Source catalog also contains the identification number of its associated DIA object (`diaObjectId`). The detections reported in the DIA Source catalog have not undergone deblending.

The DIA Source catalog contains data for 1 billion DIA sources.

#### 3.2.3. *Object*

The Object catalog contains data on all objects detected with a greater than $5\sigma$ significance in the Deep Coadds. With coadd images produced on a per-band basis, a $> 5\sigma$ detection in one or more of the bands will result in an object being included in the Object catalog. For cases where an object is detected at $> 5\sigma$ in more than one band, a crossmatching has been performed between bands to associate an object in one band with its counterpart(s) in the other bands. As such, each entry in the Object catalog associates objects across multiple bands. The objects reported in the Object catalog have also undergone deblending; in the case of blended detections, only the deblended child objects are included in the catalog. As with the Source catalog, the criterion for inclusion in the Object catalog is a $> 5\sigma$ detection in one of the Deep Coadds prior to deblending, yet the positions and fluxes of objects are reported postdeblending. Hence, it is possible for the Object catalog to contain objects whose flux-to-error ratios are less than five in all bands.

In addition to the general information mentioned above (i.e., IDs, positions, fluxes, and flags), the `Object` catalog also includes basic shape and extendedness information. While they may be included in future Data Releases, no photometric redshifts, Petrosian magnitudes (Petrosian 1976), proper motions or periodicity information are included in the DP2 object catalogs.

The `Object` catalog contains data for 790 million objects in DP2.

#### 3.2.4. *DIA Object*

The DIA Object catalog contains the astrophysical objects that the DIA sources are associated with (i.e., the DIA objects). Only non-solar-system objects are contained in the DIA Object catalog. Solar system objects are recorded in the Solar System Object catalog (§3.2.10). Once a DIA source has been detected and measured, the DIA Object and Solar System Object catalogs are searched for objects to associate it with. If no association is found, a new entry in the DIA object catalog is created and the DIA source is associated to it. An attempt has also been made to associate DIA objects across multiple bands, meaning each entry in the DIA Object catalog may contain measurements from multiple bands. Because they ultimately stem from detections in Difference Images, DIA objects are typically transient or variable. Due to this variable nature of DIA objects, the DIA Object catalog contains summary statistics of their fluxes, such as the mean and standard deviation over multiple epochs. For single epoch measurements of the properties of DIA objects, users should refer to either the Forced Source at DIA Object catalog (§3.2.6) or the DIA Source catalog for single-epoch flux measurements of DIA objects, both of which contain the DIA object identification number (`diaObjectId`) for easy look-up.

The DIA Object catalog contains data for 230 million DIA objects.

#### 3.2.5. *Forced Source*

The Forced Source catalog contains forced PSF photometry measurements performed on both Difference Images (i.e., the `psfDiffFlux` column) and Visit Images (i.e., the `psfFlux` column). Forced photometry is performed on each Visit Image and Difference Image at the positions of all the objects in the Object catalog that is covered by that image. The information in the Forced Source therefore enables the assessment of the time variability of the photometry of each object in the Object catalog. Unless there is strong justification to do otherwise, we recommend using the forced photometry measured on the Difference Images (i.e., `psfDiffFlux`) rather than the Visit Images (i.e., `psfFlux`). This is because contamination from even non-varying neighbouring sources can introduce an apparent variation `psfFlux` flux due to changes in the PSF between visits. In contrast, `psfDiffFlux` is less sensitive to contamination from neighboring sources, as non-varying sources are subtracted-out in difference images.

In addition to forced photometry PSF fluxes, a number of Boolean flags are also included in the `Forced Source catalog`.

The Forced Source catalog contains a total of 45 billion entries across the 790 million unique objects.

#### 3.2.6. *Forced Source at DIA Object*

The Forced Source at DIA Object catalog is equivalent to the Forced Source catalog, but contains forced photometry measurements performed on each Visit Image and Difference Image at the positions of all the DIA objects in the DIA Object catalog.

The Forced Source at DIA Object catalog contains a total of 24 billion entries across the 230 million unique DIA objects.

#### 3.2.7. *Isolated Star Stellar Motions*

The Isolated Star Stellar Motions catalog contains the proper motions and parallaxes of isolated stars. Rubin uses a specific definition for an isolated star. Candidate isolated stars must be point-like (i.e., unresolved) detections with a signal-to-noise ratio of between 8 and 1000, and must not be flagged as blended. All such candidates within a given tract are spatially matched using a $1''$matching radius across all visits and bands covering said tract. Should the average position of a candidate have no other candidate within $2''$, then it is considered to be an isolated star.

The Isolated Star Stellar Motions catalog provides a best-fit position (`ra, dec`), proper motion in mas $\mathrm{yr}^{-1}$ (`raPM, decPM`), and parallax in mas at a reference epoch; the `raPM` parameter already includes the cos(decl.) factor. The reference epoch in Modified Julian Date (to be avoided; see also JD) (MJD) is included as a column (`epoch`) in the catalog. An isolated star's proper motion and parallax are calculated by fitting its position across all per-visit detections, plus any detection in our calibration reference catalog (§3.2.12) that is matched to within 0.″1 of the star's average position. The catalog also contains $1\sigma$ uncertainties for each each of the five aforementioned positional parameters (five columns: raErr, decErr, raPMErr, decPMErr, parallaxErr), plus a further ten columns containing off-diagonal covariance parameters ($*$_ $*$ _Cov, e.g, ra_raPM_Cov, raPM_parallax_Cov).

All isolated stars have an entry in the Isolated Star Stellar Motion catalog, although no attempt is made to fit the time-varying positions of stars that are detected in fewer than 3 epochs and which do not have a matched counterpart in the reference catalog. In such cases the positional information, including uncertainties, are reported as NaN. Every entry in the Isolated Star Stellar Motion catalog has a unique identification number and, in cases where a match to the reference catalog has been identified, the unique identification of the matched reference source.

The Isolated Star Stellar Motions catalog contains 320 million entries.

### 3.2.8. *Object Shear*

The Object Shear catalog contains object shape measurements in five different shear steps or "counterfactual" images: one in which no shear (ns) has been applied, and four in which an artificial linear shear has been introduced to the image (1p, 1m, 2p, 2m). The artificial shears are in four different directions: 1p is along the image x-direction, 1m is along the y-direction, 2p is along the 45 degree diagonal, 2m is along the 135 degree diagonal.

Two measurement algorithms (*gauss* and *pgauss*) are run on each of the five different counterfactuals, one for measuring shapes and the other for fluxes. The first approach is to fit a single Two-dimensional (2D) gaussian jointly across the $r$-, $i$-, and $z$-band deep coadd images, with the ellipticity components in the image x-y (gauss_g1) and diagonal (gauss_g2) directions reported in the catalog. The second approach measures 2D gaussian-weighted moments on single band ($r$, $i$, $z$) images separately, calculates the weighted sum of the raw moments across those three bands (weighted by inverse-variance), then reports the ellipticity components of that weighted sum in the image x-y (pgauss_g1) and diagonal (pgauss_g2) directions. There is one row per object detected in each counterfactual, meaning a given object could result in up to five rows, depending on how many of the counterfactuals it is detected in at a significance of $> 5\sigma$.[29]

In addition to the gaussian shape measurements and their auto-covariances, the Object Shear table also contains object positional information, fluxes measured from the gaussian fits ($r$, $i$ and $z$ bands) and gaussian moments ($g$, $r$, $i$ and $z$ bands), PSF shapes reported in the same x-y and diagonal directions as objects, a range of different flags. Each entry in the Object Shear catalog also has a unique identification number (shearObjectId), that uniquely identifies a given detection in a given counterfactual. The Object Shear catalog contains 2.5 billion rows.

### 3.2.9. *Solar System Source*

The Solar System Sources catalog is derived from the Rubin observations. It contains one row per associated detection and links measured DIA Source photometry and astrometry to a predicted object in the Solar System Object catalog. As entries in the Solar System Source catalog stem from the DIA Source catalog, they have all been detected at $> 5\sigma$ significance in at least one band.

The diaSourceId can be used to identify corresponding detections in the DIA Source catalog; each diaSourceId occurs at most once in the Solar System Source catalog. The ssObjectId or designation can be used to identify corresponding entries in the object-level catalogs (i.e., Solar System Object, Minor Planet Center (MPC) Orbit, Current Identifications, and Numbered Identifications).

Each entry in the Solar System Source catalog carries the measured position and the Sorcha (Merritt et al. 2025; Holman et al. 2025) ephemeris position, the measured PSF photometry and band, the predicted apparent $V$-band magnitude, phase angle and solar elongation, heliocentric and topocentric distances, sky-plane and radial rates, and Cartesian position and velocity components. R.A. and decl. are in degrees in the ICRS frame; the observation epoch DIA_midpointMjdTai is an MJD on the International Atomic Time (TAI) time scale. State vectors are equatorial and evaluated at the light-emission (retarded) time. The ephVmag field is an approximate predicted Johnson $V$ magnitude derived from the input catalog photometric parameters. It is not a prediction in the Rubin band of the associated detection, so comparisons to the measured *ugrizy* photometry include both object color and phase-model differences.

[29] It is feasible for a marginal $5\sigma$ object in one counterfactual to be just below the $5\sigma$ detection threshold in another counterfactual.

The solar system Source catalog contains 8.1 million solar system source rows for 299819 distinct objects.

### 3.2.10. *Solar System Object*

The Solar System Object catalog is also derived from Rubin observations. It contains one row per object and summarizes its observations, per-band photometric properties, including absolute magnitudes, and selected orbit-derived quantities. The Solar System Object catalog contains 299819 entries; one entry for each distinct object. There are 17,528 objects with a single associated detection, 31,873 with at most two, 212,181 with at least ten, and 7,293 with at least 100. The main belt dominates the distance and rate distributions, but the catalog includes 35 objects observed beyond 50 astronomical unit (au), 765 source records at topocentric distances below 0.1 au, and 2,105 records with predicted rates above 1 deg day$^{-1}$. The ecliptic-latitude range of the objects, $-69.4°$ to $+13.3°$, primarily reflects the southern DP2 footprint rather than the intrinsic solar system distribution.

### 3.2.11. *MPC Orbit, Current Identifications, Numbered Identifications*

The MPC Orbit, Current Identifications, and Numbered Identifications catalogs are each auxiliary snapshots of tables in the MPC PostgreSQL database. They provide reference catalogs for orbital information and object cross-identification, retaining the upstream MPC semantics rather than being recomputed from DP2 observations. Their field definitions are documented in the MPC schema for the PostgreSQL tables replicated through the Small Bodies Node.[30] The MPC Orbits catalog records the MPC orbit solution and associated fit information for designated objects. The Current Identifications catalog identifies primary and secondary provisional designations that refer to the same object, and the Numbered Identifications catalog connects an object's permanent number to its primary provisional designation. The former includes both self-mappings ($A = A$) and links among aliases ($A = B$); the latter also carries International Astronomical Union (IAU) naming information when available. The MPC Orbits catalog available at ingestion did not contain comets. Because the predictions used for association were generated from that snapshot, no comets were available to be associated with DP2 detections and none occur in the delivered Solar System Source or Solar System Object catalogs.

In the Solar System Source, Solar System Object, and MPC Orbits catalogs, `designation` denotes the unpacked primary provisional designation of the object. This was added to the delivered NSF Mathematical and Physical Sciences directorate (MPS) Orbits copy as an end-user convenience; there it is identical, row by row, to `unpacked_primary_provisional_designation`. It therefore provides a readable and direct join key among the Solar System Source, Solar System Object, MPC Orbits catalogs, without requiring users to unpack MPC designations. For query performance, users should nevertheless prever to use integer identifiers such as `ssObjectId` to link catalogs wherever those identifiers are available, and use `designation` for readability or when an appropriate integer identifier is unavailable.



### 3.2.12. *Calibration Reference Catalog*

The Calibration Reference Catalog is used to perform astrometric and photometric calibration. It is a whole-sky catalog built specifically for LSST by combining data from multiple prior reference catalogs, since no single existing catalog had both the depth and coverage needed to calibrate LSST data. It contains only stellar sources. Full details on how the Calibration Reference Catalog was built are provided in (Ferguson et al. 2025).[31] The same reference catalog was used for calibrating DP2 as for DP1, but we prefer to include the same summary here over only providing a reference to that previous data preview.

For the *grizy* bands, the input catalogs were (in order of decreasing priority): Dark Energy Survey (DES) Year 6 Calibration Stars (Rykoff et al. 2023); Gaia B or R Photometry (Gaia) (XP) synthetic magnitudes (Gaia Collaboration et al. 2023); the Panoramic Survey Telescope and Rapid Response System (Pan-STARRS) 3PI survey (Chambers et al. 2016); Data Release 2 of the Skymapper survey (Onken et al. 2019); and Data Release 4 of the VLT Survey Telescope (VST) Asteroid Terrestrial-impact Last Alert System (ATLAS) survey (Shanks et al. 2015). For the $u$ band, the input catalogs were (in order of decreasing priority): standard stars from Sloan Digital Sky Survey (SDSS) Data Release 16 (Ahumada et al. 2020); Gaia XP Synthetic Magnitudes (Gaia Collaboration et al. 2023); and synthetic magnitudes generated using Stellar Locus Regression (SLR), which estimates the $u$-band flux from the $g$-band flux and $g$ - $r$ colors. These SLR estimates were used to boost the number of $u$-band reference sources, as other-

[30] https://docs.minorplanetcenter.net/mpc-ops-docs/data-and-services/replicated-tables-schema/

[31] In (Ferguson et al. 2025), the calibration reference catalog is referred to as "The Monster".

wise the source density from the $u$-band input catalogs is too low to be useful for the LSST.

Throughout, the Calibration Reference Catalog uses the DES bandpasses for the *grizy* bands and the SDSS bandpass for the $u$-band; color transformations derived from high-quality sources were used to convert fluxes from the various input catalogs (some of which did not use the DES or SDSS bandpasses) to the respective bandpasses. All sources from the input catalogs are matched to Gaia Data Release 3 (DR3) sources for robust astrometric information, selecting only isolated sources (i.e., no neighbors within 1″).

After collating the input catalogs and transforming the fluxes to the standard DES and SDSS bandpasses, the catalog was used to identify sources within a specific region of the sky. This process generated a set of standard columns containing positional and flux information, along with their associated uncertainties.

## 3.3. *Maps*

### 3.3.1. *Survey Property Maps*

Survey Property maps capture a number of position-dependent quantities such as observing conditions and exposure time. For each band, a given map encodes the spatial distribution of one such quantity, built up from the images that are combined to create the Deep Coadds. The construction of the Survey Property Maps is initially performed at the tract-level, then per-tract maps are combined into a single consolidated map; only these consolidated maps are provided. Maps displaying different summary statistics, such as minimum, maximum, mean, standard deviation, of the various properties are provided. Each sky position corresponds to a spatial pixel in the Hierarchical Equal-Area iso-Latitude Pixelisation (HEALPix)(Górski et al. 2005) grid.

There are forteen mag Survey Property maps in DP2. Among the quantities mapped are the total exposure time; three observation-epoch maps (earliest, mean, and latest epoch); the PSF size and shape, the latter given as separate $e^1$ and $e^2$ components; the PSF magnitude limit; sky background and noise; astrometric offsets in R.A. and decl.; and, for the wavelength-dependent atmospheric Differential Chromatic Refraction (DCR) effect, PSF distortion again split into $e^1$ and $e^2$ components.

The names of the maps follow the pattern `deepCoadd_<PROPERTY>_consolidated_map_<STATISTIC>`—so `deepCoadd_exposure_time_consolidated_map_sum`, for instance, gives the summed exposure time, in seconds, at each sky position. The maps themselves are stored in `HealSparse`[32] format, and Survey Property maps are only accessible through the Data Butler (§6.2.2).

### 3.3.2. *HiPS Maps*

Hierarchical Progressive Survey (IVOA standard) (HiPS) maps, offer an interactive way to explore seamless, multiband tiles of the sky regions covered by DP2, allowing for smooth panning and zooming. DP2 provides multiband HiPS images created by combining data from individual bands of Deep Coadd images, using an improved version (N. B. Lust et al. in preparation) of the algorithm presented in (Lupton et al. 2004). These images are false-color representations generated using various filter combinations for the red, green, and blue channels.

HiPS maps are only accessible through the HiPS viewer in the RSP Portal (§6.3) and cannot be accessed via the Data Butler (§6.2.2).

## 3.4. *Ancillary Data Products*

In addition to the science images and catalog data products described above, there are a number of ancillary data products included in DP2. While not science-level data products *per se*, their inclusion in DP2 will help to facilitate science.

### 3.4.1. *Skymap*

The DP2 skymap, `lsst_cells_v2`, is also included as a data product in DP2. The skymap data product can be used, for example, to identify the tract and patch that contains a given sky coordinate, or obtain the coordinates of the corners of a tract or patch.

A description of the `lsst_cells_v2`, in terms of the numbers and sizes of its tracts and patches, is given at the start of this section.

### 3.4.2. *Standard Passbands*

The corresponding transmission curves are provided as a DP2 data product. These datasets tabulate the full-system transmission of the six LSSTCam filters as a function of wavelength and were used as a reference for the LSSTCam DP2 photometry. The `standard_passband` dataset is keyed by band and is stored in Astropy table format.

### 3.4.3. *Object Scarlet Models*

The Object Scarlet Models dataset stores the scarlet models fit to each deblended detection in the Deep

[32] A sparse HEALPix representation that efficiently encodes data values on the celestial sphere. https://healsparse.readthedocs.io

Coadd images. Each model captures an object's spatial profile (morphology) together with its per-band flux values, as assigned by the deblender when separating it from its blended neighbours. When combined with the corresponding Object catalog and Deep Coadd image, this dataset enables the reconstruction of the deblend model to see how a specific detection was deblended. There is one Object Scarlet Models dataset per patch that has coverage in DP2, and they are only available via the Data Butler.

#### 3.4.4. *Deep Coadd Input Summary*

The Deep Coadd Input Summary table contains the visit and detector IDs of all the Visit Images that were selected and coadded to produce a given Deep Coadd image. Also included in the table are the relative weight of each Visit Image, inversely proportional to its variance, and the number of pixels it contributed to the Deep Coadd, since a given Visit Image may only partially cover a patch. There is one Deep Coadd Input Summary per tract that has coverage in DP2, and they are only available via the Data Butler.

### 3.5. *Metadata*

#### 3.5.1. *Visit*

The `Visit` catalog contains technical data for each visit collected at the time of observation, such as telescope pointing, camera rotation, airmass, exposure start and end time, and total exposure time. Some of the information contained within the `Visit` catalog is also contained in the Visit Detector catalog described below (e.g., exposure time), although the latter also includes information produced by the processing pipelines at the per-detector level, such as the PSF size and limiting magnitudes of a given Visit Image.

The Visit catalog contains 28698 entries.

#### 3.5.2. *Visit Detector*

The Visit Detector catalog contains summary information for all Visit Images in DP2. In addition to technical information, such as the on-sky coordinates of the central pixel and measured pixel scale, the Visit Detector catalog contains a range of data quality measurements, such as whole-image summary statistics for the PSF size, zero-point and $5\sigma$ depth, sky background, sky noise, and the quality of astrometric solution. It provides an efficient method to access the properties of a Visit Image without needing to access the image data. When combined with the data contained in the Visit table described above, it provides a full picture of the telescope pointing and sky conditions at the time of observation.

The Visit Detector catalog contains 5.2 million entries.

#### 3.5.3. *Visit Summary*

The Visit Summary table contains the same information as the Visit Detector catalog, but only for a single visit. There is one Visit Summary table per visit (meaning the Visit Detector table is the collation of all the Visit Summary tables), and they are only available via the Data Butler.

## 4. DATA RELEASE PROCESSING

Data Release Processing (DRP) is the data pipeline that produces the calibrated images, detection catalogs, and derived data products defined in section §3, using the LSST Science Pipelines (Rubin Observatory Science Pipelines Developers et al. 2026). DP2 was processed at the United States Data Facility (USDF) at SLAC, with pilot runs at the National Institute of Nuclear and Particle Physics (IN2P3) computing center.

The processing campaign consists of four major stages: (1) single-frame processing, (2) calibration, (3) multi-visit processing, (4) variable and transient source processing. Data quality and key metrics for verification and validation are actively monitored during each stage, and comprehensive verification and validation assessments are completed before processing the next stage.

This section describes the high-level algorithms, configurations, and verification metrics for each stage, including changes and improvements from DP1 (Vera C. Rubin Observatory Team et al. 2026).

### 4.1. *Single-Visit Processing*

This will be provided for the full DP2 release

#### 4.1.1. *Calibration & Instrument Signature Removal (ISR)*

The first task in data processing begins with a raw image. This image is in discretized analog-to-digital units (ADU, or counts) and contains detector artifacts imprinted by the physics of the detector and readout hardware. This step aims to remove any artifacts that are not of astrophysical origin and produce science-ready, single-visit exposures in floating-point electron units with complete pixel masks.

The ISR pipeline is built on a physical model of the detectors and readout electronics, in which each artifact is tied to a specific point in the photon-transfer chain and is corrected using calibrations from dedicated campaigns conducted during the day or on cloudy nights. Figure 4 shows this model of the detector and readout components and the effects each imparts on the measured image. ISR works backward through the chain, undoing these effects in the reverse of the order in which they were imprinted. Detector response and performance vary across the focal plane (see §2.2), and each detector must be calibrated individually.

A detailed description of the DP2 ISR procedures and the calibration products they consume can be found in (Broughton et al. 2026).

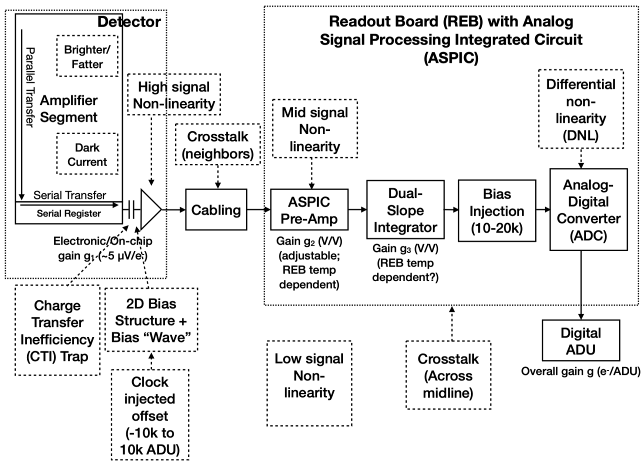


**Figure 4.** The model of the detector and Readout Electronics Board (REB) components, labeled with the effects that they impart on signal.

The ISR processing pipeline for DP2 performs the following steps, in order: (1) conversion of the raw integers to floating point, (2) Analogue-to-Digital Unit (ADU) dithering to mitigate quantization effects, (3) serial overscan subtraction, (4) saturated- and suspect-pixel masking, (5) gain (electrons/ADU) correction, (6) cross talk correction, (7) parallel overscan subtraction, (8) nonlinearity correction, (9) deferred-charge correction, (10) vendor-specific masking (E2V edge bleed; the ITL "dip," saturation sag, and edge bleed), (11) bias subtraction, (12) dark subtraction, (13) defect masking, (14) saturated pixel footprint widening, (15) brighter-fatter correction, (16) variance plane estimation, (17) flat fielding, (18) interpolation over masked pixels, (19) amplifier-offset corrections.

Compared to DP1 (Vera C. Rubin Observatory Team 2026), the ISR pipeline has been substantially upgraded, including improved sensor diagnostics and characterization, an entirely new brighter-fatter correction based on Astier & Regnault (2023) and Broughton et al. (2024), and an all-new flat-field production pipeline that corrects for gradients and mis-centering of the flat-field illumination using new calibration hardware. We generate flats using a new set of 10 LEDs spanning the full sensitivity range of the LSSTCam sensors (one in $u$, two each in $griz$, and a new 1050 nm Light-Emitting Diode (LED) shared between $y$ and $z$). For the bands covered by two LEDs, we combine separate flats for each LED (for the blue and red side of the band) into a so-called "anaglyph" flat, with a weighting that approximates a flat spectral energy distribution across the band.

Together, these upgrades allow ISR to deliver uniform, science-ready exposures across all LSSTCam science sensors for DP2.

#### 4.1.2. *Background Subtraction*

Background subtraction in `CalibrateImageTask` estimates and removes large-scale background signals from science imaging prior to source detection and measurement. Such signals include sky brightness from airglow, moonlight, scattered light, zodiacal light, and diffuse astrophysical emission. The background is estimated twice: an initial model (`psf_subtract_background`) is subtracted before PSF characterization, and a second model (`star_background`) is fit on the updated source-masked image before the final detection pass. Each post-ISR image is divided into $128 \times 128$ pixel superpixels, and the iterative $3\sigma$-clipped mean of unmasked pixels in each bin is fit with a sixth-order two-dimensional Chebyshev polynomial, evaluated at native pixel resolution via Akima spline interpolation. Before fitting `star_background`, the `DETECTED` mask plane is dilated by up to 10 pixels to suppress source-wing flux from leaking into background bins. After `star_background` subtraction, a zeroth-order pedestal correction is estimated iteratively over bin sizes starting at $32 \times 32$ pixels, doubling each step until the cumulative pedestal level changes by less than 5% or 0.5 counts between steps. The `star_background` model and pedestal corrections together constitute the full background model.

Compared to DP1 (NSF-DOE Vera C. Rubin Observatory 2025), several changes were made to background subtraction for DP2. The inline re-estimation performed after each detection pass in DP1 (`reEstimateBackground = True`) has been replaced by the dedicated `star_background` subtask. In addition to source mask dilation and a pedestal correction, other new features for DP2 include adaptive threshold determination and diffraction spike masking. The per-step illumination correction used in DP1 has been disabled for DP2. Following background subtraction, the median and standard deviation of all unmasked pixels and of *sky source* fluxes are recorded in the task metadata as diagnostics of background subtraction quality.

#### 4.1.3. *Source Detection*

Pending

## 4.2. *Calibration*

### 4.2.1. *PSF Modeling*

PSF modeling in DP2 uses two distinct approaches, as described in the (Rubin Observatory Science Pipelines Developers et al. 2026) and it is similar to what was originally described in DP1. The first model, which is faster but less precise, is fit with PSFEx (Bertin 2011), while the final model, which is slower but more accurate, is fit with Piff (Jarvis et al. 2021). Broadly speaking, PSFEx and Piff both represent the PSF on a pixel basis and interpolate its parameters across a single CCD using two-dimensional polynomials. However, their technical implementations differ (see Bertin et al. and Jarvis et al. 2021 for more details). Previous comparisons have shown that Piff outperforms PSFEx by delivering an unbiased reconstruction of the PSF size (see Zuntz et al. 2018 and Jarvis et al. 2021), can work in sky coordinates and can include a color dependency, and supports alternative models and interpolation schemes (see Schutt et al. 2025 for more details).

In the initial PSF modeling pass, PSFEx is run with a pixel basis and a second-order polynomial interpolation on each CCD. For the final modeling, Piff uses its PixelGrid with a fourth-order polynomial interpolation per CCD—except in the $u$-band, where star counts are too low to constrain a fourth-order fit, so a second-order polynomial is used instead. Both are working in pixel coordinates.

### 4.2.2. *Astrometric Calibration*

The final astrometric calibration of the data follows the same method as for DP1, but with some key additions. As described in (NSF-DOE Vera C. Rubin Observatory 2025), a joint calibration is performed on all visits in a given band overlapping a given area. By associating all the isolated point sources shared among the overlapping visits, we can refine the astrometric solution beyond the level achieved in single-frame processing using only a reference catalog (Saunders 2024). The full details of the method are described in Saunders (2024).

Unlike in DP1, the area we use for individual fits of the astrometric solution is an order-3 healpix, instead of a tract. This is done to reduce repetitive calculations, since the small size of tracts compared to the LSSTCam field of view means that neighboring tracts often share identical overlapping visits, which would lead to doing many identical fits of the astrometric calibration.

As in DP1, the astrometric model consists of a static map from pixel space to an intermediate frame, which is followed by a per-visit map from the intermediate frame to the plane tangent to the telescope boresight, and finally a deterministic mapping from the tangent plane to the sky. This is done by the `gbdes` package (Bernstein et al. 2017), as described in Saunders (2024). The proper motion and parallax of the objects used in the calibration are fit as part of the solution, but we do not yet include differential chromatic refraction or lateral color. These will be included in future Data Releases during Rubin Operations.

The main difference from the DP1 astrometric calibration is an additional step that models the effect of atmospheric turbulence on the astrometry. The per-visit part of the astrometric model already described captures large-scale variation between visits, which includes some atmospheric effects. However, higher order oscillations are visible in the astrometric residuals, and these can be attributed to atmospheric turbulence (Saunders 2025). As described in Fortino et al. (2021) and Léget et al. (2021), the wave-like astrometric effects caused by turbulence can be modeled by a Gaussian process. In DP2, we do this by taking the astrometric residuals remaining after the `gbdes` fit, grouping them by visit, then modeling them with the Gaussian Process package `treegp`. We then use this Gaussian Processes model to predict the effects of turbulence on a grid. Finally, this grid is turned into a two-dimensional spline that can use the AST component `SplineMap` so that it can be concatenated with the `gbdes` model. By including this model for the higher-order atmospheric turbulence, we achieve a significant improvement in the astrometric calculation.

A corollary product of the `gbdes` astrometric calibration is the catalog of proper motions and parallaxes of the sources used in the calibration. Since the `gbdes` fit is done per-band, and then improved by the turbulence fit, we perform an additional final fit of the proper motion and parallax, using the final astrometric model and all available bands. The resulting catalogs are included in the DP2 Data Release. These catalogs only include isolated point sources, but future Data Releases will be extended to include proper motion and parallax for all objects.

### 4.2.3. *Photometric Calibration*

Photometric calibration of the DP2 dataset is based on the Forward Global Calibration Method (FGCM) (Burke et al. 2018) adapted for the LSST Science Pipelines (Aihara et al. 2022; Fagrelius & Rykoff 2025). We used the FGCM to calibrate the full DP2 dataset with a forward model that uses a parameterized model of the atmosphere as a function of airmass along with a model of the instrument throughput as a function of wavelength. The FGCM process typically begins with measurements of the instrumental throughput, including the mirrors, filters, and detectors. For the LSST-

Cam filter throughputs, we start with the filter scans provided by the manufacturer, Matereon. These were taken in a grid with 159 positions, at 5 incidence angles from 13.8 to 23 degrees, at a wavelength resolution of 1 nm. We performed an averaging procedure to get the approximate appropriate throughput for the LSSTCam beam as it intersects the filter at the central location of each detector. For the LSSTCam detector quantum efficiency (QE) curves, we start with per-raft measurements of the QE performed at Brookhaven National Laboratory (BNL) and SLAC at 10 to 25 wavelength steps. These preliminary scans were then normalized relative to each other at 750 nm. However, there are still some per-raft systematic differences seen, particularly in the u and y bands. In the future we will use the in-dome tunable laser system and the CBP for improved filter and detector throughputs.

The FGCM atmosphere model, provided by MODTRAN (Berk et al. 1999), was used to generate a look-up table for atmospheric throughput as a function of zenith distance at Cerro Pachón. This model accounts for absorption and scattering by molecular constituents of the atmosphere, including O2 and O3; absorption by water vapor; and Mie scattering by airborne aerosol particulates. Nightly variations in the atmosphere are modeled by minimizing the variance in repeated observations of stars with an Signal to Noise Ratio (SNR) greater than 10, measured using "compensated aperture fluxes." These fluxes include a local background subtraction to mitigate the impact of background offsets. The model fitting process incorporates all six bands, *ugrizy*, but does not include any gray (achromatic) terms, except for a linear assumption of mirror reflectance degradation and dust accumulation on lenses. Fortunately, there is no evidence of any signficant loss of throughput over the 250 days of the DP2 dataset. As an additional constraint on the fit, we use a subset of stars from the Calibration Reference Catalog[33], primarily to constrain the system's overall throughput and establish the "absolute" calibration.

## 4.3. *Multi-Visit Processing*

### 4.3.1. *Input Data Selection*

As described in §2.2, detectors 0, 20, 27, 65, 123, 161, 168, 188 are excluded from the `deep_coadd` images as they are strongly affected by vignetting. Only exposures with a seeing better than 1.″7 FWHM and transparency better than 0.75 are included in the `deep_coadd` images. Exposures with poor PSF modeling, identified using internal diagnostics, are also excluded from the `deep_coadd` images to prevent contamination of the coadds with unreliable PSF estimates.

### 4.3.2. *Coaddition*

The single Visit Images that pass the selection criteria above are warped onto the coadd pixel grid using Lanczos3 kernel and scaled to have units of `nJy` (zeropoint of 31.4). Artifacts are identified and masked in a manner similar to DP1 before coaddition. We apply the artifact rejection procedure described in (AlSayyad 2018), which identifies and masks features such as satellite trails, optical ghosts, and cosmic rays. It operates on a time series of PSF-matched images resampled onto a coadd pixel grid ("warps") and leverages their temporal behavior to distinguish persistent astrophysical sources from transient artifacts. Artifact rejection uses both direct (where no PSF matching is performed) and PSF-matched warps, homogenized to a standard PSF of 1.″8 FWHM, broadly consistent with the 1.″7 FWHM seeing threshold used in data screening. A sigma-clipped mean of the PSF-matched warps serves as a static sky model, against which individual warps are differenced to identify significant positive and negative residuals. Candidate artifact regions are classified as transient if they appear in fewer than a small percentage of the total number of exposures, with the threshold based on the number of visits Note: note released as data product in DP2

In DP2, the coadded images are cell-based (§3.1.3), using the `lsst_cells_v2` skymap. Coadds are built as weighted averages of the input warps. Each detector is assigned a weight that is the inverse of the mean of the variance reported in the Visit Summary dataset (§3.5.3). Following (Armstrong et al. 2024) for each cell in the coadd, a detector is excluded if its geometry does not overlap the cell completely. Furthermore, a detector is excluded if more than 50% of its pixels are masked as artifacts, unless doing so would result in an empty cell in which case, this condition is ignored. Finally, pixels that did not get any contribution from the input visits are interpolated from neighboring pixels.

A key difference between cell-based coadds and coadds in DP1 is in the PSF representation. Single visit PSFs are evaluated at a point corresponding to the center of each cell and warped into a 35x35 pixel postage stamp, which ensures that at least 99.5% of the flux is captured. These postage stamps are then weighted by the same weights used for the coadd and averaged. The PSF within the cell is approximated as spatially constant. The pixels that have artifacts from any of the input visits have `INEXACT_PSF` flag set to denote that

[33] Referred to as "The Monster" (§3.2.12; see Ferguson et al. 2025

the PSF of those pixels may differ from that attached to the cell. The metadata about which visits contribute to the cell, what fraction of pixels are unmasked, PSF summary statistics are stored.

#### 4.3.3. *Detection, Deblending, & Measurement*

Coadd detection and measurement proceeds with a similar order of steps as in DP1; however, numerous changes have been made to individual algorithms. The main change is that algorithms run on cell-based coadds, which have fewer PSF discontinuities within cells. For this release, though, none of the measurement algorithms take cell boundaries into account, and those that require a PSF model use the PSF model at the centroid of the object (and therefore of the cell that the centroid is within) for all cells.

First, peaks are detected in single-band patch images. These peaks are then merged across bands to produce a multiband peak catalog. This catalog is used as an input for deblending, which outputs deblended (but not deconvolved) images for each object. All object catalog measurements are performed on these deblended images.

The basic methodology for detection is largely unchanged from the description in Bosch et al. (2018), which we summarize here. To detect peaks, each image is convolved by a circular Gaussian approximation to the measured PSF to make a significance image. A footprint is a contiguous region of pixels above the detection threshold, which is set to detect isolated point sources above $5\sigma$ significance. Footprints can have multiple peaks, which are defined as pixels with values higher than all 8 neighbours (local maxima within a 3x3 box). Note that even for an isolated object, the significance of a detection is not precisely the same as the signal-to-noise from subsequent PSF model fluxes.

Pixels within footprints are marked with the `DETECTED` bit mask. The remaining pixels are used for background estimation and subtraction, excluding significantly negative-valued pixels which are likely to be residual artifacts. The algorithm also attempts to randomly place 100 "sky objects" within the noise-dominated pixels, reducing the number in crowded fields with low numbers of sky-dominated pixels if necessary. These objects do not appear in the final object catalog but are used to characterize the sky variance in intermediate steps.

As explained in Aihara et al. (2019), the process of convolving the image by an approximation of the PSF reduces the variance in the image, and therefore the pixel significance threshold must be adjusted to achieve the desired detection threshold. This factor is computed by measuring the standard deviation of the PSF model SNR of the sky objects.

Once detection has run in each available band, footprints and their peaks are merged in a fixed priority order, *irzygu*, which is roughly in order of expected SNR for typical objects. New peaks are only added if they are at least $1\farcs0$ further than an existing peak, and existing peaks are marked as detected in multiple bands if a new peak is found within $0\farcs3$; peaks within this range are ignored. After merging, probable spurious peaks detected in only one band are removed.

Deblending is performed by the `ScarletLite` python package. `ScarletLite` is an optimized implementation of the scarlet (Melchior et al. 2018) deblending algorithm with the assumption that the input images are aligned to the same pixel grid at a fixed resolution, as is the case in coadds. Prior to deblending, a new addition in DP2 is a per-band partial deconvolution of the input coadds to a common $\sigma = 0.8$ pixel integrated circular Gaussian PSF. This allows us to sub-divide large parent blends (footprints) into multiple smaller blends that are only connected by the wings of the PSF. These smaller sub-blends and the initialization of galaxy morphologies from the partially deconvolved images improves numerical convergence and performance, both in the goodness-of-fit of the final models and in memory usage and runtime. Users will notice that this has improved our coverage[34] in crowded fields from DP1; however, in the densest fields, an entire patch can still be a single blend and must be skipped due to excessive memory requirements.

Object measurements are conducted in three modes: independent per-band measurements, forced measurements in each band, and multiband measurements. As detailed in Bosch et al. (2018), measurement first replaces all pixels in footprints with noise, so that algorithms that require pixel values beyond footprint boundaries can perform measurements without contamination from neighbors (whether they are part of the same blend or not). During measurement, isolated objects have their original pixel values inserted back into the image, while blended objects use the scarlet lite models as templates to re-distribute the original image values. This preserves the flux from the original image where objects have no overlap and dampens the noise elsewhere, since the flux is reapportioned to multiple objects.

The single-band measurement plugin system remains unchanged from DP1 and largely follows the description in Bosch et al. (2018). Measurements are made in every available band, and a reference band is defined from the same priority order as detection merging (*irzygu*) for a second round of forced measurements. These use the

[34] the fraction of each field with usable measurements

shape and position from the reference for consistency. Most algorithms have their free measurements reported in the final object table, but forced values are generally recommended.

A variety of flux measurements are provided, including aperture fluxes, Kron (1980), and Gaussian-Aperture-and-PSF (Gaussian Aperture and PSF (GAaP) Kuijken 2008; Kannawadi 2025) fluxes for consistent galaxy colors. Forward PSF model and Composite model (CModel) magnitudes (Abazajian et al. 2004; Bosch et al. 2018) are provided for each band.

MultiProFit (Taranu 2025) performs multiband forward modelling. Beginning in DP2, an exponential profile is fit first and used to initialize a subsequent Sérsic model (Sérsic 1963; Sersic 1968) fit, which improves convergence and reduces the prevalence of Sérsic fits with unrealistically large effective radii. Both exponential and Sérsic model fits have free centroids, which serve as alternative, multiband astrometry for objects. The Object table now includes uncertainties on the fitted ellipse parameters for both models (the x- and y-axis effective radii and correlation parameter $\rho$). These will be converted to uncertainties on more commonly used parameters (major/minor effective radii and axis ratios) for Data Release 1 (DR1).

Some additional derived parameters are included in the object catalog, mainly for object classification. "extendedness" is a binary classifier, set to 1 if the PSF to CModel flux ratio is below 0.985. "sizeExtendedness" is a quasi-probabilistic morphological classifier reported as $1 - \exp(-0.5 * ((T_{\mathrm{HSM}} - T_{\mathrm{PSF}})/T_{\mathrm{PSF}}^{0.5})^2)$, where $T_{\mathrm{HSM}}$ and $T_{\mathrm{PSF}}$ are the sum of the squares of the x- and y-axis moments for the object from the HSM algorithm and for the PSF model, respectively. As an alternative, DP2 adds `model_extendedness`, which is a continuous, heuristic classifier that uses both the exponential model effective radius and the ratio of the Sérsic model to PSF model flux. It is available for each band and separately for the sum of the *griz*-band fluxes.

As for DP1, photometric redshifts are not computed or included in the object table but are planned as a supplementary user-generated product.

Further details on the performance of these algorithms are found in §5.

#### 4.3.4. *Metadetection and ShearObject*

The ShearObject catalog using the multiband metadetection algorithm (Sheldon et al. 2020) is new to DP2. It creates a separate catalog of detections that do not correspond to entries in the Object catalog. Each 150x150 cell in *griz* bands is extracted with a 50 pixel border to make 250x250 pixel images, that are then passed to metadetect. Internally, metadetect makes counterfactual images by deconvolving by the original PSF, shearing the image and reconvolving to a new PSF using the `fitgauss` reconvolution algorithm. Metadetect coadds *riz* bands to form a coadd for detection, and measures shear using the `gauss` algorithm by fitting a 2D elliptical Gaussian to $r$, $i$, $z$ coadds. Additionally, for these detections, it also reports force-photometered fluxes in $g$, $r$, $i$ and $z$ bands for these detections measured using the `pgauss` algorithm. There is no deblending performed prior to these measurements.

### 4.4. *Variable & Transient Source Processing*

#### 4.4.1. *Difference Image Analysis (DIA)*

DP2 difference imaging analysis did not change significantly relative to DP1. DP2 DIA uses the decorrelated Alard and Lupton image differencing algorithm (Alard & Lupton 1998; Reiss & Lupton 2016). Both positive and negative `DIASource`s are detected in the difference image at $5\sigma$. We fit a dipole model to sources containing both positive and negative peaks, and infer trail parameters from the second moments of the PSF. After rejecting likely artifacts (§4.4.2), we associate `DIASource`s into `DIAObject`s with a $1\farcs0$ matching radius.

#### 4.4.2. *Artifact Rejection*

DIA analysis is prone to false positives. We employ several layers of artifact rejection, although we err on the side of completeness rather than purity in Data Release Processing.

As in DP1, prior to writing the `DIASource` catalog, we remove sources with pixel flags artifacts, non-astrophysical trail lengths, and unphysically negative direct fluxes. In DP2 we also reject sources that land on the modeled positions of diffraction spikes and apply morphological cosmic ray rejection on both the direct and difference images. Further, we flag but do not remove multiple `DIASources` that occur in a straight line with `glint_trail`, indicating that they may be due to rotating orbital debris (Karpov & Peloton 2022).

We developed a machine-learned reliability score to assess whether `DIASources` were astrophysical. This model is a convolutional neural network that was trained on template, science, and difference-image postage stamps around `DIASources` using a combination of simulated fake injections and human-labeled detections. (Acero-Cuellar et al. 2026) provides an overview of the model architecture, training, and performance of the real-bogus model. The v0.1 model used in DP1 was again used during DP2 processing. No cuts were applied on reliability for the DP2 `DIASource` catalogs, so users interested in higher purity should apply a minimum reliability threshold.

#### 4.4.3. *Light Curves*

As with DP1, multiepoch forced photometry is performed on both the Visit Images and the difference images. For users interested in light curve construction, we recommend using the forced photometry performed on the difference-images (the `psfDiffFlux` column of the ForcedSource table), since this measurement isolates the time-variable component of the flux and is unaffected by contamination from non-varying nearby sources. Forced photometry carried out directly on the Visit Images, by contrast, may pick up flux from blended or nearby static objects, and the severity of this contamination depends on the seeing. The centroids adopted for this multiepoch forced photometry step are drawn either from coadd-based object positions or from the DIAObjects, which are constructed from the DIASources associated from detections in the difference images.

### 4.5. *Solar System*

#### 4.5.1. *Source Association*

The solar system association input was a 2025 March 13 snapshot of the MPC Orbit database[35], limited to objects whose observational arcs exceeded two days. Associations were made independently for each detector and visit using positions predicted by Sorcha (Merritt et al. 2025; Holman et al. 2025) at the visit epoch from this orbit catalog. Predicted objects were first restricted to the detector footprint. Measured DIA Source (§3.2.2) coordinates were represented as unit vectors on the celestial sphere and stored in a $k$-d tree. For each predicted position, the algorithm identified the nearest detection within $1''$. Candidate pairs were considered in increasing separation, accepting a pair only if neither member had already been assigned. Thus, this is a positional, one-to-one match; it uses neither brightness, color, morphology, nor rate and does not calculate a probabilistic association score. The delivered `diaDistanceRank` is set to one for every row and should not be used as a measure of match ambiguity. Chance associations remain possible because the match does not use brightness or detection reliability, and are more likely where the DIA source density is high. Users constructing a high-purity sample should consider detection reliability, angular and cross-track residuals, consistency with the predicted magnitude, and the uncertainty of the input orbit. Detection-level reliability information requires reference to the DIA Source catalog (§3.2.2).

#### 4.5.2. *Photometry and Phase Curves*

The solar system Object photometric fields were derived independently in each of the *ugrizy* bands from PSF fluxes in the associated detections. For a flux $f$ in nJy, the conversion used was $m_{\rm AB} = 31.4 - 2.5\log_{10} f$, with magnitude uncertainty $1.085736\,\sigma_f/f$. Non-positive fluxes therefore yield non-finite magnitudes and were removed before fitting. Apparent magnitudes were reduced to unit heliocentric and topocentric distance using $m(1,1,\alpha) = m - 5\log_{10}(r\Delta)$, where $r$ and $\Delta$ are in au.

The reduced magnitudes were described with the $H, G_{12}$ phase function of Muinonen et al. (2010). The DP2 processing fixed $G_{12} = 0.5$ and fit only the absolute magnitude $H$, since the short observing interval and limited phase-angle coverage do not generally support a stable two-parameter fit. Accordingly, valid per-band `G12` values are all 0.5, while `G12Err` and the $H$–$G_{12}$ covariance are undefined. A 0.05 mag uncertainty floor was added in quadrature to the measurement uncertainties to reduce sensitivity to unmodeled scatter, including rotational variability. After removing non-finite magnitudes, an initial fit with a robust loss was used to reject residuals beyond $10\sigma$, followed by a conventional least-squares fit to the retained observations. The per-band `nObsUsed` field records the retained count and should be compared with `nObs` when assessing clipping or invalid photometry.

Valid $H$ estimates were obtained for 35,268, 196,015, 224,888, 231,890, 203,341, and 90,789 objects in *ugrizy*, respectively. Among objects with sufficient observations for a fit, failures range from 53 in $y$ to 1,034 in $g$, or less than approximately 0.5% in every band. Depending on band, 1–5% of the eligible measurements were removed as invalid or clipped. There are 25,706 objects with no fitted $H$ in any band; many are represented by only one detection and hence have insufficient information for a fit. The median fitted $H$ values are 18.7, 18.2, 17.5, 17.3, 17.2, and 16.6 mag in *ugrizy*. These statistics are dominated by main-belt objects and are also shaped by the DP2 footprint, cadence, phase-angle coverage, and band-dependent detection selection; they should not be interpreted as intrinsic population size distributions.

## 5. PERFORMANCE CHARACTERIZATION AND KNOWN ISSUES

For information regarding the performance and known issues of the data contained in the Early DP2 release, we refer to the Data Quality, Image Artifacts, and Known Issues sections of the Release Documentation at https:

[35] Small bodies discovered after 2025 March 13—including those found in Rubin data—are therefore absent from the solar system Source catalog §3.2.9

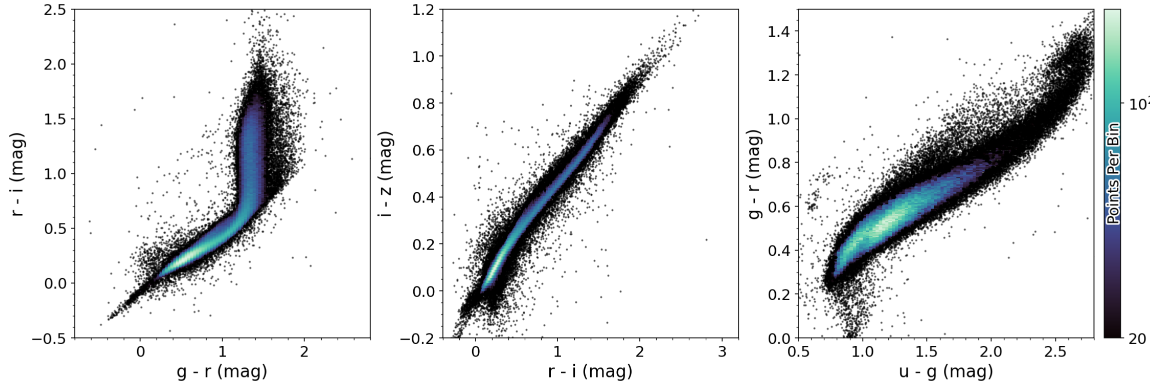


**Figure 5.** Blah

//dp2.lsst.io. A detailed analysis of the validation of DP2 will be provided at the time of its release (expected in the Oct–Dec 2026 range).

### 5.1. *Stellar Photometry*

Since most of the DP2 footprint area was observed multiple times, we can use the multiple photometric measurements of the same star to assess the consistency, and hence reliability, of the per-visit photometric calibration. This photometric repeatability is measured by calculating the standard deviation of multiple measurements of the same unresolved object (the vast majority of which will be stars) across all visits that observed it. We only use bright (i.e., SNR> 200), isolated, unresolved objects (i.e., bright, isolated stars) to perform this measurement. To make this calculation computationally feasible, however, photometric repeatability is measured at the tract-level: the same set of stars are identified in each visit that cover a given tract. For each star, the standard deviation of the multiple photometric measurements across all visits is measured. A star must have at least three measurements in order for the standard deviation to be calculated. We then calculate the median standard deviation over all stars in a tract to obtain a single value for the photometric repeatability within that tract. This is referred to as "PA1" (Ivezić & The LSST Science Collaboration 2018). Following recalibration using the FGCM fits, the photometric repeatability is excellent: across all tracts with coverage, the median PA1 values for the *ugrizy* bands are 9.1, 7.1, 6.8, 5.9, 5.8, and 6.2, mmag respectively.

We turn to stellar loci in *ugriz* to assess the quality of the photometry of Deep Coadds; these are shown in Figure **??**. Unresolved objects are selected using the Object table's extendedness parameter. This parameter is a measure of the difference between the PSF flux and the CModel flux; a large value (

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) inidicates a substantial difference betweem the CModel fit and the PSF model, which is taken as implying that the object is resolved. The narrowness

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of the stellar loci are testament to the quality of the photometry measured in the Deep Coadds.

Neither the repeatability nor the stellar locus tests described above are sensitive to errors in the absolute flux scale, since both are differential measures (i.e., between visits or between bands) in which any common offsets will cancel. As a simple check of absolute calibration, we compare DP2 photometry to Gaia

confirm which Gaia data product/release

, finding a median offset of

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and a scatter of

quantify, per band

in the

bands

bands. This comparison should, however, be interpreted with caution. Gaia's own photometry carries calibration systematics of its own – most notably chromaticity-related biases in the BP/RP spectra used to synthesize magnitudes in the LSST passbands, which are known to be more significant for faint and red sources – so agreement (or disagreement) with Gaia cannot be uncritically attributed to DP2. The comparison is also restricted to the relatively bright, mostly unsaturated stars in the magnitude range where Gaia and LSST overlap, and is therefore not representative of the full DP2 depth. Furthermore, converting Gaia spectra into the LSST passbands is itself a synthetic-photometry exercise, so this test is better described as a comparison between two calibration systems than as a validation against an independent physical standard. A more thorough assessment of DP2's absolute photometric calibration – including a proper treatment of these systematics, alongside complementary tests such as recovery of the fluxes of sources injected into the images at known brightness – is left to a dedicated future paper on photometric performance.

## 6. RUBIN SCIENCE PLATFORM

The RSP (Jurić et al. 2019) is a powerful, cloud-based environment for scientific research and analysis of petascale astronomical survey data and serves as the primary interface for scientists to access, visualize, and conduct next-to-the-data analysis of Rubin and LSST data. Built around a "bring the compute to the data" principle, it eliminates the need for users to download massive datasets, instead providing Rubin data rights holders with secure, single-sign-on access to data and services via three distinct user-facing Aspects: a *Portal*, which facilitate interactive exploration of the data; a JupyterLab-based *Notebook* environment for data analysis using Python; and an extensive set of *APIs* that enable programmatic access to both data and services.

The three Aspects are designed to be fully integrated, enabling seamless workflows across the RSP. The RSP is supported by a number of backend services, including databases, files, and batch computing. Support for collaborative work through shared workspaces is also included in the RSP. The data products described in §3 are accessible via all three Aspects, and the system facilitates operations such as starting a query in one Aspect and retrieving its results in another. Figure 6 shows the RSP landing page in the Google cloud.

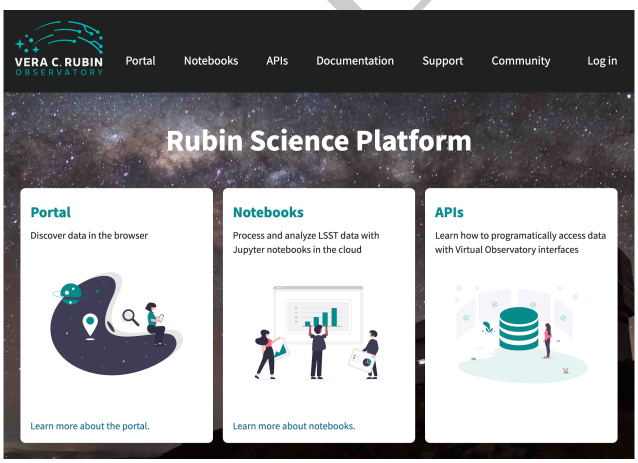


**Figure 6.** The RSP landing page at https://data.lsst.cloud/ showing the three user-facing Aspects as well as links to documentation and support information.

A preview of the RSP was first launched on Google Cloud in 2022, operating under a shared-risk model to support DP0 (O'Mullane et al. 2024b). DP1 built on that early preview to bring major upgrades including the hybrid on-premises USDF/cloud deployment, a broader set of International Virtual Observatory Alliance (IVOA) API services (SIAv2, SODA cutouts, HiPS, WebDAV, alongside the existing TAP), and a fully integrated, read-only Data Butler (S3 store + AlloyDB registry) across all three RSP Aspects.

The size (3.5 TB) and complexity of DP1 did not, on their own, necessitate the full functionality of the RSP for analysis; a traditional "download the data" model would still have sufficed. However, it was important to provide users the same experience and interfaces they will rely on for later, larger releases, and to get them started in this new paradigm on a small, manageable dataset. DP2, by contrast, marks a new regime, with its size of 59.15 TB outside the range of the traditional download model. The experience users gained working with DP1 through the RSP interfaces can now be applied directly to DP2. The release of DP2 builds on this evolving platform to bring additional upgrades and new functionality, including expanded user query-history and query-status capabilities, context-aware in-application documentation (e.g., contextual help linked directly from Portal table columns), and support for the ADQL spatial `INTERSECTS()` operation.

The RSP remains under active development, with incremental improvements being rolled out as they mature. During the Rubin Early Science Phase, the RSP will continue to operate under a shared-risk model. This section outlines the RSP functionality available at the time of the DP2 release and provides an overview of planned future capabilities.

### 6.1. *Rubin US Data Access Center*

The Rubin US DAC utilizes a novel hybrid architecture that combines on-premises infrastructure at the USDF at SLAC with flexible, scalable resources in the Google cloud (O'Mullane et al. 2024a). User-facing services are deployed in the cloud to support dynamic scaling in response to user demand and to simplify the provisioning and management of large numbers of science user accounts, whereas the majority of the static data products described in §3 are stored on premises at the USDF, benefiting from cost-effective mass storage and close integration with Rubin's data processing infrastructure, also located at the USDF. For all imaging data except coadds, the Data Butler (§6.2.2) provides the interface between cloud-based users and the on-premises data; for catalog data, a cloud-based TAP client (§6.2.1) submits queries to the on-premises Qserv database cluster (§6.5.1) and retrieves the results. Deep Coadd images are the exception to this hybrid model: because there is one coadd per sky region per band, a relatively small and size-bounded dataset, they are hosted directly in Google Cloud to benefit from cloud-native access.

DP2 is being delivered in two phases because full end-to-end integration of the hybrid RSP for images was not yet complete by the time data processing (§ 4) finished; the hybrid model supporting Qserv (§6.5.1) catalog access over TAP, by contrast, is already in production for EDP2. Rubin therefore adopted a phased release, given that neither coadds nor catalogs depend on the image-side hybrid RSP integration readiness. Early DP2 (EDP2), comprising the deep coadd images and all catalogs, was released on 2026-07-27, while the full DP2 release, adding the visit-level data products once hybrid image integration is demonstrated, is expected in the Oct–Dec 2026 range.

### 6.2. *Application Programming Interface Aspect*

The Application Programming Interface (API) Aspect provides a comprehensive set of user-facing interfaces for programmatic access to the DP1 data products, through both IVOA-compliant services and the Rubin Data Butler. IVOA services enable standard queries and integration with existing tools, while the Butler facilitates advanced data processing with the LSST Science Pipelines.

Since the DP1 release, additional services have been added as the RSP moves toward full functionality, but some planned services remain unavailable, and certain data products are only accessible via the Butler. This section provides an overview of the available IVOA services and Butler access.

#### 6.2.1. *IVOA Services*

Rubin has adopted a Virtual Observatory (VO)-first design philosophy, prioritizing compliance with IVOA standard interfaces to foster interoperability, standardization, and collaboration. In cases where standardized protocols have yet to be established, additional services have been introduced to complement these efforts. This approach ensures that the RSP can be seamlessly integrated with community-standard tools such as Tool for OPerations on Catalogues And Tables (TOPCAT) (Taylor 2011) and Aladin (Bonnarel et al. 2000; Boch & Fernique 2014; Baumann et al. 2022), as well as libraries such as PyVO (Graham et al. 2014).

The user-facing APIs are also used internally within the RSP, creating a unified design that ensures consistent and reproducible workflows across all three Aspects. This reduces code duplication, simplifies maintenance, and ensures all users, both internal and external, access data in the same way. For example, an Astronomical Data Query Language (IVOA standard) (ADQL) query on the `Object` catalog via TAP yields identical results whether run from the Portal, Notebook, or an external client.

Since the DP1 release, the RSP has continued to progress towards a full-featured platform, with a number of new services and capabilities added in support of DP2. The following IVOA services are available at the time of the DP2 release.

- **TAP service**: A TAP service (Dowler et al. 2019) enables queries of catalog data via the IVOA-standard ADQL, a dialect of SQL92 with spherical geometry extensions. The main TAP service for DP2 runs on the Rubin-developed Qserv database (§ 6.5), which hosts the core science catalogs described in §3.2, as well as the Visit and Visit Detector catalogs. It also provides image metadata in the IVOA `ObsCore` format via the standard `ivoa.ObsCore` table, making it an "ObsTAP" service (Louys et al. 2017). The TAP service is based on the Canadian Astronomy Data Centre (CADC)'s open-source Java TAP implementation[36], modified for the exact query language accepted by Qserv. Since DP1, ADQL support has been expanded to include the spatial `INTERSECTS()` operation, and temporary-table uploads now provide multi-object query capabilities across both Qserv- and Postgres-backed catalogs. It currently supports a large subset of ADQL, with limitations documented in the Data Release materials (see §7) and exposed via the TAP capabilities endpoint where possible.

  The TAP service provides metadata annotations consistent with the standard, including table and column descriptions, indications of foreign-key relationships between tables, and column metadata such as units and IVOA unified content descriptors (UCDs).

- **Image access services**: Rubin image access services are compliant with IVOA Simple Image Access Protocol, version 2 (SIAv2; Jenness et al. 2024; Dowler et al. 2015) for discovering and accessing astronomical images based on metadata. SIAv2 is a REpresentational State Transfer (REST)-based protocol designed for the discovery and retrieval of image data. It allows, for instance, querying all images in a given band over a defined sky region and time period.

[36] https://github.com/opencadc/tap

Users identify an image or observation of interest and query the service. The result set includes metadata about the image, such as the sky position, time, or band, and a data-access URL, which includes an IVOA Identifier uniquely identifying the dataset (Jenness & Dubois-Felsmann 2026), allowing the dataset to be retrieved or a cutout requested via Server-side Operations for Data Access (IVOA standard) (SODA) (SODA; Bonnarel et al. 2017).

- **Image Cutout Service**: The Rubin Cutout Service (Allbery 2023, 2024) is based on the IVOA SODA standard (Bonnarel et al. 2017). Users submit requests specifying sky coordinates and the cutout size as the radius from the coordinates, and the service performs the operation on the full image and returns a result set. As with DP1, the Cutout Service for DP2 remains a single-cutout service, where $N$ cutout requests require $N$ independent synchronous calls; a bulk Cutout Service, (Jenness et al. 2026) is now planned for DR1.

- **HiPS Data Service**: An authenticated HiPS (Fernique et al. 2017) Data Service for seamless pan-and-zoom access to large-scale coadds. It supports fast interactive progressive image exploration at a range of resolutions.

- **Web Distributed Authoring and Versioning (WebDav)**: A WebDav service is provided to enable users to remotely manage, edit, and organize files and directories on the RSP as if they were local files on their own computer. This is especially useful for local development.

Beyond these IVOA services, DP2 introduces expanded user query-history and query-status capabilities, initial context-aware, in-application documentation (e.g., contextual help linked directly from Portal table columns).

Looking ahead to DR1, planned RSP functionality includes bulk cutout services (both for individual targets across a range of epochs and for lists of multiple targets), a PSF retrieval service, an on-demand forced-photometry service, data product recreation services, parallel and batch computing, support for persistent user-created databases and their federation with project-provided databases, support for collaborative work, and full LSSTCam focal-plane visualizations. Access to GPUs and the ability to bring additional user-funded compute resources to the RSP are under consideration for the post-DR1 period.

### 6.2.2. *Data Butler*

The Rubin Data Butler (Jenness et al. 2022; Lust et al. 2023), is a high-level interface designed to facilitate seamless access to data for both users and software systems. This includes managing storage formats, physical locations, data staging, and database mappings. A Butler repository contains two components.

- The *Data Store*. A physical storage system for datasets, e.g., a Portable Operating System Interface (POSIX) file system or (Amazon) Simple Storage Service (S3) object store.

- The *Registry*. A Structured Query Language (SQL)-compatible database that stores metadata about the datasets in the Data Store.

In the context of the Butler, a dataset refers to a unique data product, such as an image, catalog or map, generated by the observatory or processing pipelines. Datasets belong to one of the various types of data products, described in §3. The Butler ensures that each dataset is uniquely identifiable by a combination of three pieces of information: a data coordinate, a dataset type, and a run collection.

For a DP2 deep coadd on a patch of sky, the dimensions would be the tract, patch and band, specified as a dataId, e.g., `dataId={'tract': 11399, 'patch': 16, 'band': 'g'}` and is associated with the `deep_coadd` DatasetType. In contrast, the `visit_table` and `visit_detector_table` catalogs described in §3 have only the instrument dimension, since each is a single table covering all visits (or all visit-detector combinations) for a given instrument; retrieving either therefore requires only `dataId={'instrument': 'LSSTCam'}`. The `visit_summary` catalog, by contrast, is keyed on band, instrument, `day_obs`, `physical_filter`, and visit, e.g., `dataId={'instrument': 'LSSTCam', 'visit': 2025071100074, 'band': 'g', 'physical_filter': 'g_6', 'day_obs': 20250711}`, and contains one row per detector for that visit.

As for DP1, the EDP2 Butler repository containing the catalogs and coadds described in §3 is hosted in the Google Cloud using an S3-compatible store for datasets and AlloyDB, a PostgreSQL-compatible database for the Registry.

### 6.2.3. *Remote Programmatic Access*

The Rubin RSP API can be accessed remotely by data rights holders by creating a user security token. This token can then be used as a bearer token for API calls to the RSP services. This capability is especially useful for remote data analysis using tools such as TOPCAT, and for enabling third-party systems, such as community Alert brokers, to access Rubin data. Additionally, it supports remote development, allowing for more flexible workflows and integration with external systems.

### 6.3. *Portal Aspect*

The Portal Aspect provides an interactive web-based environment for exploratory data discovery, filtering, querying, and visualization of both image and catalog data, without requiring programming expertise. It enables users to access and analyze large datasets via tools for catalog queries, image browsing, time-series inspection, and crossmatching.

The Portal is built on Firefly (Wu et al. 2019), a web application framework developed by the Infrared Processing and Analysis Center. Firefly provides interactive capabilities such as customizable table views, image overlays, multipanel visualizations, and synchronized displays linking catalog and image data.

Designed to support both exploratory data access and detailed scientific investigation, the Portal delivers an intuitive user experience, allowing users to visually analyze data while retaining access to underlying metadata and query controls.

### 6.4. *Notebook Aspect*

The Notebook Aspect provides an interactive, web-based environment built on Jupyter Notebooks, enabling users to write and execute Python code directly on Rubin and LSST data without downloading it locally. It offers programmatic access to Rubin and LSST data products, allowing users to query and retrieve datasets, manipulate and display images, compute derived properties, plot results, and reprocess data using the LSST Science Pipelines (§4). The environment comes preinstalled with the Pipelines and a broad set of widely used astronomical software tools, supporting immediate and flexible data analysis.

### 6.5. *Databases*

The user-facing Aspects of the RSP are supported by several backend databases that store catalog data products, image metadata, and other derived datasets. The schema for DP2 and other Rubin databases are available online at https://sdm-schemas.lsst.io/dp2.html.

#### 6.5.1. *Qserv*

The final 10 yr LSST catalog is expected to reach 15 PB and contain measurements for billions of stars and galaxies across trillions of detections. To support efficient storage, querying, and analysis of this dataset, Rubin Observatory developed Qserv (Wang et al. 2011; Mueller et al. 2025), a scalable, parallel, distributed SQL database system. Qserv partitions data over approximately equal-area regions of the celestial sphere, replicates data to ensure resilience and high availability, and uses shared scanning to reduce overall I/O load. It also supports a package of scientific user-defined functions (SciSQL[37]) simplifying complex queries involving spherical geometry, statistics, and photometry. Qserv is built on robust production-quality components, including MariaDB[38] and XRootD[39]. Qserv runs at the USDF and user access to catalog data is via the TAP service (§6.2.1). This enables catalog-based analysis through both the RSP Portal and Notebook Aspects.

While DP1 was small enough that Qserv's full capabilities were not strictly required, DP2, at 59.15 TB, is substantially larger and begins to exercise Qserv at a scale meaningfully closer to its intended operating regime, though still well short of the petabyte-scale catalogs expected for the 10 yr LSST survey.

Qserv is open-source and available on GitHub at https://github.com/lsst/qserv.

#### 6.5.2. *OpSim*

DP2 provides an opsim-format SQLite database of visit-level observational metadata for the visits included in DP2, downloadable from https://s3df.slac.stanford.edu/data/rubin/sim-data/prelsst/dp2_visits.db, with additional documentation available at https://survey-strategy.lsst.io/progress/prelsst/prelsst_dp2.html#visit-database. Like the Visit catalog descibed in §3.5.1, this database contains visit-level entries, but it has no equivalent of the Visit Detector catalog described in §3.5.2 and so provides no per-detector information. Its quantities (e.g., PSF FWHM, 5$\sigma$ point-source depth) are drawn from the real-time Rapid Analysis QuickLook processing, () performed on the summit rather than from DRP processing. Values may therefore differ somewhat from the corresponding DRP-derived quantities in the Butler `visit_table`. The database also includes some quantities not otherwise available from DRP outputs, such as AOS Zernike coefficients, as well as predicted quantities useful for comparison against simulated survey performance.

## 7. SUPPORT FOR COMMUNITY SCIENCE

Rubin Observatory's model to support its large, diverse science community to access and analyze DP2 fo-

[37] https://smonkewitz.github.io/scisql/
[38] https://www.mariadb.org/
[39] https://xrootd.org/

cuses on providing documentation, tutorials, and live virtual learning opportunities alongside an open platform for issue reporting and resolution that enables crowd-sourced solutions. These support resources are all available via the main hub of Rubin's "For Scientists" website[40].

The Data Release documentation for DP2[41] provides an overview of the LSSTCam observations, data quality summary plots, and descriptions of the contents and access methods for all catalog and image datasets. All tutorials (NSF-DOE Vera C. Rubin Observatory 2021) that demonstrate DP2 access via the Rubin Science Platform are delivered in the Data Release documentation, and executable versions of the Jupyter notebook tutorials are also available from a drop-down menu within the Notebook Aspect.

Early DP2 was released on the first day of the annual Rubin Community Workshop, July 27 2026. A dedicated "Get started with DP2" session was added to the Wednesday agenda. Four DP2-onboarding virtual sessions were run in the first week of August, a mix of presentations, tutorials, and time for issue resolution.

The Rubin Community Forum[42] is the open platform for user issue reporting and resolution. Rubin staff monitor all new topics posted in the Support category and ensure that all requests for help get a response and a solution. The open nature of the platform also enables peer-to-peer support that avoids knowledge bottlenecks and hastens the time to resolution.

## 8. SUMMARY AND FUTURE RELEASES

DP2 provides the first release of science-grade data products from the LSSTCam focal plane. As for DP1, DP2 is part of Rubin's Early Science Program (Guy et al. 2026) and carries a 2 yr proprietary period during which access is restricted to LSST data rights holders via the cloud-based RSP.

DP2 is delivered in two phases: Early DP2 (EDP2), which provides deep coadded images, all science catalogs, and ancillary data products including throughput curves, survey property maps, and calibration products; and DP2, planned for late 2026, which adds single-epoch visit-level and difference-image data products to complete the release. The coadded images and catalogs delivered in EDP2 will not be reprocessed for DP2; the data are unchanged, and DP2 supersedes and fully encompasses EDP2 by adding the remaining products to the same release. Analyses begun on EDP2 products require no modification.

DP2 represents a substantial advance relative to DP1: the full LSSTCam focal plane replaces the single-raft LSSTComCam, all six *ugrizy* filters are available in place of three, and the sky area is larger by approximately two orders of magnitude. Nevertheless, DP2 reflects the state of an observatory still in the process of optimization: delivered image quality had not yet consistently reached design specification, template coverage is incomplete across the full sky, and several pipeline components remain under active development. Users are encouraged to consult the release documentation at https://dp2.lsst.io and the known issues register maintained there before beginning scientific analyses.

The body of science-grade LSSTCam data acquired between the end of the DP2 observing window on 2026-01-06 and the formal start of the LSST on 2026-06-29 will be processed and made available to LSST data rights holders through the same RSP services. The format and timing of this release are still under discussion; planning will be finalised once DP2 has been delivered.

The first LSST Data Release, DR1, will be derived from a reprocessing of approximately the first year of LSST survey data and is currently expected in mid-2028; its precise definition remains under consideration. DR1 will mark a qualitative advance beyond DP2: observations collected under the community-developed baseline survey strategy (Rubin's Survey Cadence Optimization Committee et al. 2022, 2023, 2025, 2026) will yield substantially broader and more uniform sky coverage across the full WFD footprint of 18000 square degrees, a balanced six-band filter distribution, and improved delivered image quality as AOS and dome thermal management continue to mature (NSF-DOE Vera C. Rubin Observatory 2026b). Photometric and astrometric calibration will benefit from a full year of all-sky coverage, difference-imaging templates will span a far greater fraction of the footprint and processing will draw on a more mature version of the LSST Science Pipelines refined through DP2 experience and community engagement. Subsequent LSST Data Releases are planned at an approximately yearly cadence, each a complete reprocessing of all survey data accumulated to that point, with progressively deeper coadds and improved calibration. Plans will be updated through revisions to Guy et al. (2026) and on the Rubin community forum at https://community.lsst.org.

[40] https://rubinobservatory.org/for-scientists

[41] https://dp2.lsst.io

[42] https://community.lsst.org/

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*Software:* Rubin Data Butler (Jenness et al. 2022), LSST Science Pipelines (Rubin Observatory Science Pipelines Developers et al. 2026), LSST Feature Based Scheduler v3.0 (Jones 2026; Yoachim et al. 2024; Naghib et al. 2019), Astropy (Astropy Collaboration et al. 2013, 2018, 2022), PIFF (Jarvis et al. 2021), GBDES (Bernstein 2022), MultiProFit (Taranu 2025), Qserv (Wang et al. 2011; Mueller et al. 2025),

## APPENDIX

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Glossary

**2D:** Two-dimensional.

**ADQL:** Astronomical Data Query Language (IVOA standard).

**ADU:** Analogue-to-Digital Unit.

**airmass:** The pathlength of light from an astrophysical source through the Earth's atmosphere. It is given approximately by sec z, where z is the angular distance from the zenith (the point directly overhead, where airmass = 1.0) to the source.

**algorithm:** A computational implementation of a calculation or some method of processing.

**AOS:** Active Optics System.

**API:** Application Programming Interface.

**astrometry:** In astronomy, the sub-discipline of astrometry concerns precision measurement of positions (at a reference epoch), and real and apparent motions of astrophysical objects. Real motion means 3-D motions of the object with respect to an inertial reference frame; apparent motions are an artifact of the motion of the Earth. Astrometry per se is sometimes confused with the act of determining a World Coordinate System (WCS), which is a functional characterization of the mapping from pixels in an image or spectrum to world coordinate such as (RA, Dec) or wavelength.

**ATLAS:** Asteroid Terrestrial-impact Last Alert System.

**au:** astronomical unit.

**AURA:** Association of Universities for Research in Astronomy.

**AuxTel:** LSST's 1.2-meter Auxiliary Telescope will measure atmospheric transmission and will be used to calibrate LSST images.

**background:** In an image, the background consists of contributions from the sky (e.g., clouds or scattered moonlight), and from the telescope and camera optics, which must be distinguished from the astrophysical background. The sky and instrumental backgrounds are characterized and removed by the LSST processing software using a low-order spatial function whose coefficients are recorded in the image metadata.

**BNL:** Brookhaven National Laboratory.

**Butler:** A middleware component for persisting and retrieving image datasets (raw or processed), calibration reference data, and catalogs.

**CADC:** Canadian Astronomy Data Centre.

**calibration:** The process of translating signals produced by a measuring instrument such as a telescope and camera into physical units such as flux, which are used for scientific analysis. Calibration removes most of the contributions to the signal from environmental and instrumental factors, such that only the astronomical component remains.

**camera:** An imaging device mounted at a telescope focal plane, composed of optics, a shutter, a set of filters, and one or more sensors arranged in a focal plane array.

**CBP:** Collimated Beam Projector.

**CCD:** Charge-Coupled Device.

**cloud:** A visible mass of condensed water vapor floating in the atmosphere, typically high above the ground or in interstellar space acting as the birthplace for stars. Also a way of computing (on other peoples computers leveraging their services and availability).

**configuration:** A task-specific set of configuration parameters, also called a ‘config’. The config is read-only; once a task is constructed, the same configuration will be used to process all data. This makes the data processing more predictable: it does not depend on the order in which items of data are processed. This is distinct from arguments or options, which are allowed to vary from one task invocation to the next.

**COSMOS:** Cosmic Evolution Survey.

**CTI:** Charge Transfer Inefficiency.

**DAC:** Data Access Center.

**DAQ:** Data Acquisition System.

**Data Management System:** The computing infrastructure, middleware, and applications that process, store, and enable information extraction from the LSST dataset; the DMS will process peta-scale data volume, convert raw images into a faithful representation of the universe, and archive the results in a useful form. The infrastructure layer consists of the computing, storage, networking hardware, and system software. The middleware layer handles distributed processing, data access, user interface, and system operations services. The applications layer includes the data pipelines and the science data archives’ products and services.

**Data Release Processing:** Deprecated term; see Data Release Production.

**DCR:** Differential Chromatic Refraction.

**DDF:** Deep Drilling Field.

**deblend:** Deblending is the act of inferring the intensity profiles of two or more overlapping sources from a single footprint within an image. Source footprints may overlap in crowded fields, or where the astrophysical phenomena intrinsically overlap (e.g., a supernova embedded in an external galaxy), or by spatial co-incidence (e.g., an asteroid passing in front of a star). Deblending may make use of a priori information from images (e.g., deep CoAdds or visit images obtained in good seeing), from catalogs, or from models. A ’deblend’ is commonly referred to in terms of ’parent’ (total) and ’child’ (component) objects.

**DECam:** Dark Energy Camera.

**deg:** degree; unit of angle.

**DES:** Dark Energy Survey.

**DIA:** Difference Image Analysis.

**DMS:** Data Management Subsystem.

**DOE:** Department of Energy.

**DP0:** Data Preview 0.

**DP1:** Data Preview 1.

**DP2:** Data Preview 2.

**DR1:** Data Release 1.

**DR3:** Data Release 3.

**DRP:** Data Release Processing.

**E2V:** Teledyne-e2v semiconductor company.

**ECDFS:** Extended Chandra Deep Field-South Survey.

**EDFS:** Euclid Deep Field South.

**EDP2:** Early Data Preview 2.

**EPO:** Education and Public Outreach.

**epoch:** Sky coordinate reference frame, e.g., J2000. Alternatively refers to a single observation (usually photometric, can be multi-band) of a variable source.

**exponential profile:** The radial distribution of flux of an astronomical source that is characterized: I(r)=I0exp(.68(r/re)) The normalization 1.68 is chosen so that the model radius is a half-light radius. An 2-dimensional elliptical version of this profile is fit to every detected source.

**FBS:** Feature-Based Scheduler.

**FGCM:** Forward Global Calibration Method.

**Firefly:** A framework of software components written by IPAC for building web-based user interfaces to astronomical archives, through which data may be searched and retrieved, and viewed as FITS images, catalogs, and/or plots. Firefly tools will be integrated into the Science Platform.

**flux:** Shorthand for radiative flux, it is a measure of the transport of radiant energy per unit area per unit time. In astronomy this is usually expressed in cgs units: erg/cm$2$/s.

**footprint:** See 'source footprint', 'instrumental footprint', or 'survey footprint', 'Footprint' is a Python class representing a source footprint.

**forced photometry:** A measurement of the photometric properties of a source, or expected source, with one or more parameters held fixed. Most often this means fixing the location of the center of the brightness profile (which may be known or predicted in advance), and measuring other properties such as total brightness, shape, and orientation. Forced photometry will be done for all Objects in the Data Release Production.

**FWHM:** Full Width at Half-Maximum.

**GAaP:** Gaussian Aperture and PSF.

**Gaia:** a space observatory of the European Space Agency, launched in 2013 and expected to operate until 2025. The spacecraft is designed for astrometry: measuring the positions, distances and motions of stars with unprecedented precision.

**HEALPix:** Hierarchical Equal-Area iso-Latitude Pixelisation.

**HiPS:** Hierarchical Progressive Survey (IVOA standard).

**HSM:** Shape measurement algorithm from Hirata & Seljak (2003) and Mandelbaum et al. (2005).

**IAU:** International Astronomical Union.

**ISR:** Instrument Signal Removal.

**ITL:** Imaging Technology Laboratory (UA).

**IVOA:** International Virtual Observatory Alliance.

**L1:** Lens 1.

**L2:** Lens 2.

**L3:** Lens 3.

**LED:** Light-Emitting Diode.

**LSST:** Legacy Survey of Space and Time.

**LSSTCam:** LSST Science Camera.

**LSSTComCam:** Rubin Commissioning Camera.

**LWS:** Light-Wind Screen.

**M1:** Primary Mirror.

**M1M3:** Single piece of glass for Primary Mirror/Tertiary Mirror.

**M2:** Secondary Mirror.

**metadata:** General term for data about data, e.g., attributes of astronomical objects (e.g. images, sources, astroObjects, etc.) that are characteristics of the objects themselves, and facilitate the organization, preservation, and query of data sets (e.g., a FITS header contains metadata).

**MJD:** Modified Julian Date (to be avoided; see also JD).

**MPC:** Minor Planet Center.

**MPS:** NSF Mathematical and Physical Sciences directorate.

**MultiProFit:** The MultiProFit astronomical source modelling code, including the multiprofit Python package and its dependencies, and the meas_extensions_multiprofit package for pipeline tasks.

**NIST:** National Institute of Standards and Technology (USA).

**NSF:** National Science Foundation.

**Object:** In LSST nomenclature this refers to an astronomical object, such as a star, galaxy, or other physical entity. E.g., comets, asteroids are also Objects but typically called a Moving Object or a Solar System Object (SSObject). One of the DRP data products is a table of Objects detected by LSST which can be static, or change brightness or position with time.

**Pan-STARRS:** Panoramic Survey Telescope and Rapid Response System.

**passband:** The window of wavelength or the energy range admitted by an optical system; specifically the transmission as a function of wavelength or energy. Typically the passband is limited by a filter. The width of the passband may be characterized in a variety of ways, including the width of the half-power points of the transmission curve, or by the equivalent width of a filter with 100% transmission within the passband, and zero elsewhere.

**patch:** An quadrilateral sub-region of a sky tract, with a size in pixels chosen to fit easily into memory on desktop computers.

**pipeline:** A configured sequence of software tasks (Stages) to process data and generate data products. Example: Association Pipeline.

**POSIX:** Portable Operating System Interface.

**PPDB:** Prompt Products DataBase.

**PSF:** Point Spread Function.

**PTC:** Photon Transfer Curve.

**PVI:** Processed Visit Image.

**QE:** quantum efficiency.

**Qserv:** LSST's distributed parallel database. This database system is used for collecting, storing, and serving LSST Data Release Catalogs and Project metadata, and is part of the Software Stack.

**REB:** Readout Electronics Board.

**REST:** REpresentational State Transfer.

**RSP:** Rubin Science Platform.

**S3:** (Amazon) Simple Storage Service.

**scarlet lite:** The lightweight version of the scarlet algorithm, optimized for multiband deblending on pixel-aligned grids. Includes the scarlet_lite and meas_extensions_scarlet Python packages.

**schema:** The definition of the metadata and linkages between datasets and metadata entities in a collection of data or archive.

**SDSS:** Sloan Digital Sky Survey.

**shape:** In reference to a Source or Object, the shape is a functional characterization of its spatial intensity distribution, and the integral of the shape is the flux. Shape characterizations are a data product in the DIASource, DIAObject, Source, and Object catalogs.

**ShearObject:** An object with measurements using the metadetection algorithm, or the catalog thereof.

**Simonyi Survey Telescope:** The telescope at the Rubin Observatory that will perform the LSST (this refers to all physical components: the mirror, the mount assembly, etc.).

**SLAC:** SLAC National Accelerator Laboratory.

**SLR:** Stellar Locus Regression.

**SNR:** Signal to Noise Ratio.

**SODA:** Server-side Operations for Data Access (IVOA standard).

**software:** The programs and other operating information used by a computer.

**Solar System Object:** A solar system object is an astrophysical object that is identified as part of the Solar System: planets and their satellites, asteroids, comets, etc. This class of object had historically been referred to within the LSST Project as Moving Objects.

**Source:** A single detection of an astrophysical object in an image, the characteristics for which are stored in the Source Catalog of the DRP database. The association of Sources that are non-moving lead to Objects; the association of moving Sources leads to Solar System Objects. (Note that in non-LSST usage 'source' is often used for what LSST calls an Object.).

**SQL:** Structured Query Language.

**SV:** Science Validation.

**TAI:** International Atomic Time.

**TAP:** Table Access Protocol (IVOA standard).

**TBD:** To Be Defined (Determined).

**TOPCAT:** Tool for OPerations on Catalogues And Tables.

**tract:** A portion of sky, a spherical convex polygon, within the LSST all-sky tessellation (sky map). Each tract is subdivided into sky patches.

**transient:** A transient source is one that has been detected on a difference image, but has not been associated with either an astronomical object or a solar system body.

**USDF:** United States Data Facility.

**Visit:** A sequence of one or more consecutive exposures at a given position, orientation, and filter within the LSST cadence. See Standard Visit, Alternate Standard Visit, and Non-Standard Visit.

**VO:** Virtual Observatory.

**VST:** VLT Survey Telescope.

**WebDav:** Web Distributed Authoring and Versioning.

**WFD:** Wide-Fast-Deep.

**XP:** B or R Photometry (Gaia).